



Noise, Regression Dilution Bias, and Solar-Wind/Magnetosphere Coupling Studies

Joseph E. Borovsky*

Space Science Institute, Boulder, CO, United States

Using numerical experiments, the effects of noise in the solar-wind and magnetospheric data on fits to the data are examined. In particular, the impact of noise amplitude on the functional forms of best-fit solar-wind driver functions is explored. The presence of noise (measurement error) will make it difficult to use solar wind and magnetosphere data to uncover (or confirm) the formula that describes the physics of the driving of the magnetosphere.

Keywords: magnetosphere, solar wind, geomagnetic activity, geomagnetic indices, solar wind magnetosphere coupling, space weather

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*Correspondence:

Joseph E. Borovsky
jborovsky@spacescience.org

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INTRODUCTION

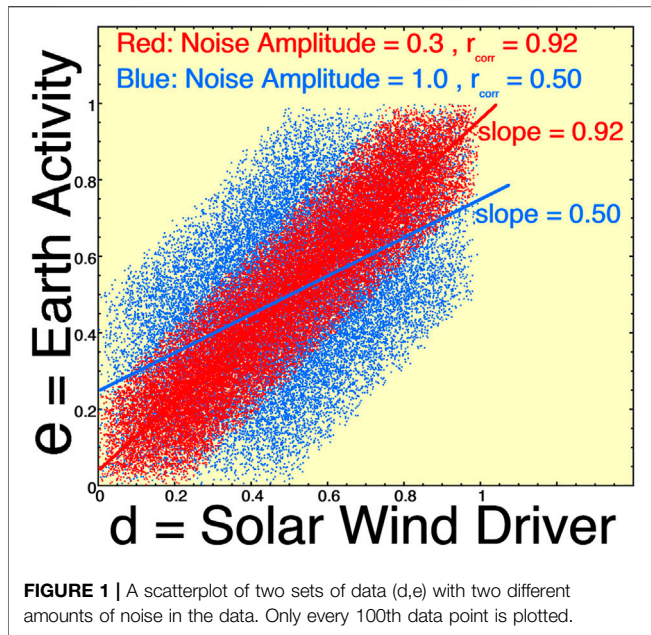
Solar-wind/magnetosphere coupling is often studied by examining “driver functions” created from multiple solar-wind variables and testing how well the driver functions do in statistically describing the time-dependent activity of the Earth’s magnetosphere-ionosphere system, with that activity typically measured with a single geomagnetic index. Often the goodness of the driver function is measured by the magnitude of the Pearson linear-correlation coefficient between the time-dependent solar-wind driver function and the time-dependent geomagnetic index. Correlation coefficients of 0.5–0.8 are typical.

Associated with the linear correlation, a least-squares linear-regression fit to the geomagnetic-index values as a function of the driver-function values is often made. In a plot (for example, **Figure 1**) of the geomagnetic index (vertical) versus the solar-wind driver function (horizontal), the least squares fit is based on minimizing the vertical errors from a line on the plot. In a sense, this least-squares linear-regression fit is the best fit for predicting the value of the geomagnetic index (vertical) knowing the value of the driver (horizontal).

In this report, artificial data sets are used to explore the effects of noise in the data for the study of solar-wind/magnetosphere coupling. For simplicity and clarity, the artificial data sets employed will not involve time lags as the actual solar-wind and magnetospheric data do.

REGRESSION DILUTION BIAS

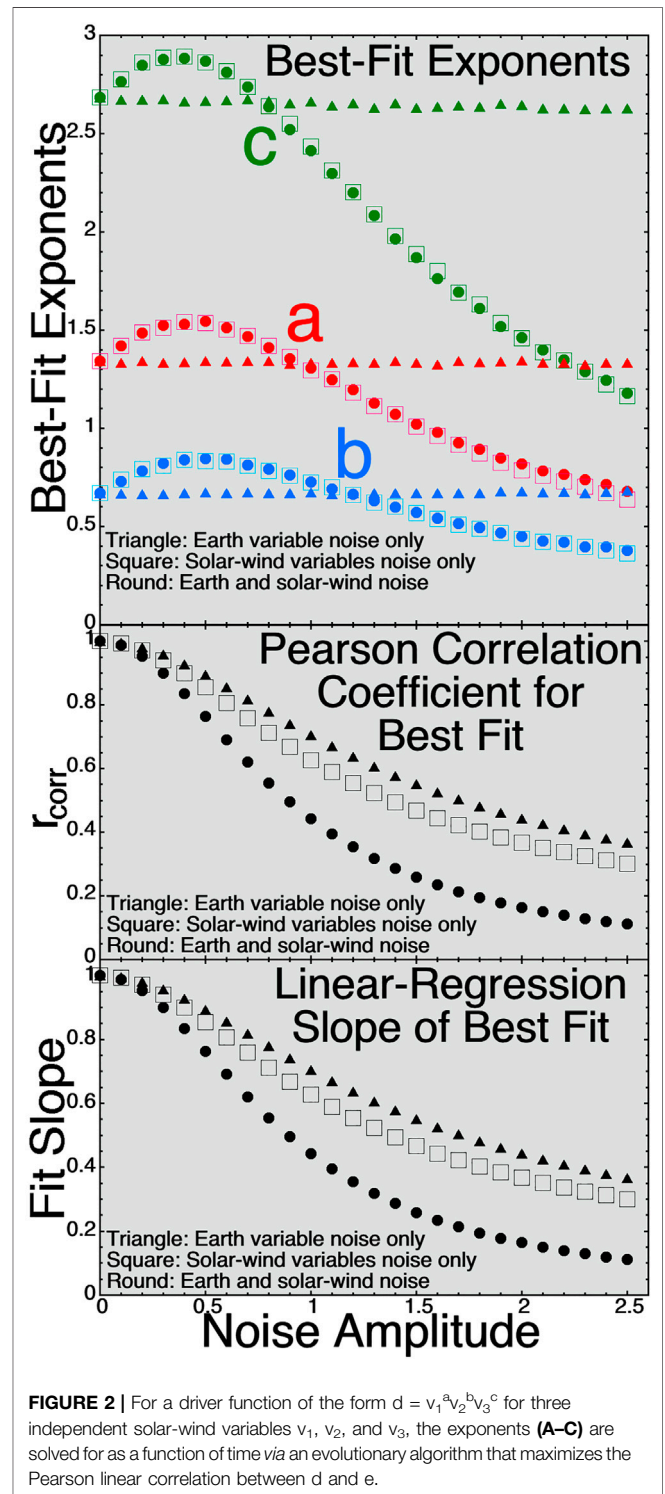
Data that is imperfectly correlated leads to a phenomenon denoted as “regression dilution bias” (e.g., Liu, 1988; Hutcheon et al., 2010) or as “attenuation by errors” (e.g., Spearman, 1904; Bock and Petersen, 1975). Basically, the smaller the Pearson correlation coefficient r_{corr} , the shallower the slope of the linear regression fit: that is the systematic “bias”. Hence, the larger the noise in the data, the lower the correlation coefficient, and the shallower the slope of the linear-regression fit. Additionally for data points x versus y , the linear-regression fit formula obtained for $y(x)$ (y fitted as a function of x) differs from the fit formula for $x(y)$ (x fitted as a function of y).



In some sense a better fit to the data is a “major-axis linear-regression fit” (Riggs et al., 1978; Warton et al., 2006), also known as a “total least squares fit” (Golub and Van Loan, 1980) or a “Gaussian fit” (Borovsky et al., 1998): this fit minimizes the perpendicular distances to the line rather than just minimizing the vertical distances to the line. If you were to “eyeball” a scatterplot and draw a line through the group of points, your line would approximate the major axis fit and would have a slope steeper than the mathematical linear-regression least-squares fit.

Figure 1 displays some of these concepts with artificial data. Data points e (Earth activity, vertical) are plotted as a function of d (solar wind driver, horizontal). The data sets each are comprised of 300,000 points (d,e), although only every 100th point is plotted. The core data set (d_o, e_o) is not plotted, but it is created as follows. d_o (solar-wind driver) is a box-car distribution of random numbers between 0 and 1. Then e_o (Earth) is created as $e_o = d_o$. If e_o were to be plotted as a function of d_o , all points would lie on the line $e = d$, the slope of the linear-regression fit would be 1.0, and the Pearson correlation coefficient would be $r_{corr} = 1.0$. The red points in **Figure 1** are created by adding noise (boxcar random numbers) to both d_o and e_o where the boxcar noise values go from -0.15 to +0.15. The d and e distributions ($d = d_o + \text{noise}$ and $e = e_o + \text{noise}$) are then “standardized” so that they go from values of 0 to values of 1. Similarly the blue points in **Figure 1** are created by adding larger-amplitude noise to the d_o and e_o points, where the boxcar noise goes from -0.25 to +0.25, and the distributions are “standardized” after the noise is added. Least-square linear regression fits are performed and plotted as the two lines: a red line for the red points and a blue line for the blue points. For the red points the fit slope is 0.92 and for the more-noisy blue points the fit slope is 0.5. Recall that the “true answer” if there was no noise in the data would be a slope of 1.0. As noted in the scatterplot of **Figure 1**, with increasing noise the Pearson linear correlation coefficient r_{corr} is reduced.

If the physics of the solar wind driving the magnetosphere is $e = d$ as described by the $e_o = d_o$ points, then noise in the variables in **Figure 1** is yielding systematically different formulas for the driving: $e = 0.92d$ and $e = 0.5d$. With increasing inaccuracy of the data, the interpretation of the fit formulas is that the solar-wind driving of the earth is weaker than it should be: the increase in Earth activity associated with an increase in driving is lessened.



EFFECT OF NOISE ON A BEST-FIT FORMULA

The solar-wind driver functions are mathematical combinations of solar wind variables. The functional forms used are most often multiplicative combinations of solar-wind variables with non-unity exponents on some of the variables (cf. Table 1 of Baker, 1986, Table 1 of Newell et al., 2007, Table 1 of Balikhin et al., 2010, Table 1 of Borovsky, 2013), or they can be linear combinations of solar wind variables (Borovsky and Denton, 2018; Borovsky, 2021), or they can be time integrals of solar-wind variables (Borovsky, 2017). We don't know the "correct" functional form of the solar-wind driver function for the Earth, so we often look for the solar-wind function that gives the best correlation with geomagnetic indices (e.g., Newell et al., 2007; Borovsky, 2014; McPherron et al., 2015). Let's ask whether noise in the data changes those combinations, i.e., whether noise changes the functional form of a best-fit solar-wind formula to describe the Earth activity.

For a mathematical gedanken experiment, let's suppose we know how the driving works and can describe it with a solar wind formula. **Figure 2** explores how noise in the solar-wind-magnetosphere data can change the functional form of best-fit solar-wind driver functions. As in **Figure 1** a core data set (d_0, e_0) is created, where here the solar-wind driver function d_0 is constructed from three independent solar-wind variables v_{10} , v_{20} , and v_{30} represented by three sets of 100,000 random numbers. The driver function will be taken to have a functional form like the Newell driver (Newell et al., 2007) $d_0 = v_{10}^{4/3} v_{20}^{2/3} v_{30}^{8/3}$. (The Newell function is $v_{sw}^{4/3} B_{sw}^{2/3} \sin^{8/3}(\theta_{clock}/2)$.) In the reference data set (d_0, e_0) of 100,000 point pairs the Earth reaction is taken to be $e_0 = d_0$. Let's assume d_0 is the driver function that describes the physics of the driving and e_0 is the real reaction of the Earth to d_0 . As was the case in **Figure 1**, noise will be added to d_0 and e_0 to make various noisy data sets (d, e) . The added noise are random numbers. The "noise amplitude" is the standard deviation of the noise-number distribution divided by the standard deviation of the variable to which the noise is added. The noise will be added in three different manners: 1) noise added only to e_0 (vertical noise on the e-versus-d scatter plot), 2) noise added only to v_{10} , v_{20} , and v_{30} (horizontal noise on the e-versus-d scatter plot), and 3) noise added to both the vertical and the horizontal. For each noisy data set v_1 , v_2 , v_3 , and e the following calculation is made. An evolutionary algorithm (genetic algorithm) (cf. Borovsky, 2017; Borovsky, 2020a) is run to solve for the three exponents a , b , and c such that the Pearson correlation between the driver function $d = v_1^a v_2^b v_3^c$ and the earth function e is maximum. The algorithm randomly changes the values of a , b , and c : if a random change produces a driver $d = v_1^a v_2^b v_3^c$ with a larger correlation coefficient r_{corr} , then the change is accepted: if the random change produces a lower correlation coefficient, then the change is rejected and the formula is reverted back to the pre-change form. The algorithm evolves a , b , and c to a local maximum in r_{corr} . There is no guarantee that there is only one local maximum, but whenever the algorithm has been run with drastically different initial values of a , b , and c it evolves to the same final set of a , b , and c values. In the top panel of **Figure 2** the values of a , b , and c that give the maximum correlation are plotted as a function of the amplitude of

the noise added to v_{10} , v_{20} , v_{30} , and e_0 . The three shapes of the points correspond to the three separate ways the noise was added. In the middle panel of **Figure 2** the maximum correlation coefficient r_{corr} for that amount of noise obtained by the algorithm between $d = v_1^a v_2^b v_3^c$ and e for the best-fit a , b , and c values is plotted. As expected, the correlation coefficient r_{corr} decreases with increasing noise amplitude. Note however in the top panel that the best-fit values of a , b , and c vary with the noise amplitude if there is noise in the solar-wind variables (round and hollow-square points). Recall that the answer in the absence of noise was $a = 4/3$, $b = 2/3$, and $c = 8/3$ such that $d_0 = v_{10}^{4/3} v_{20}^{2/3} v_{30}^{8/3}$. Let's call d_0 the formula describing the physics of the solar-wind driving the magnetosphere. As **Figure 2** demonstrates, with noise (which there always is in measurements of the solar wind for the real magnetosphere) the data yields a different formula from the one that describes the "physics". The changing of the values of a , b , and c in the driver formula $d = v_1^a v_2^b v_3^c$ is what this report considers as a changing of the functional form of the driver function caused by noise.

In the bottom panel of **Figure 2** the slopes of linear-regression fits to the e values as functions of the best-fit d values are plotted as functions of the noise amplitude. (Both d and e are standardized here, with mean values of 0 and standard deviations of 1.) The slope values in the bottom panel track the correlation coefficients in the middle panel, commensurate with the regression-dilution-bias effect. I.e., for the linear best fit of e by $v_1^a v_2^b v_3^c$, the coefficient in front of $v_1^a v_2^b v_3^c$ decreases with increasing noise.

Note that if there is vertical-only noise on the Earth measure (geomagnetic index) e but not in the solar wind, the coefficients obtained would not change with noise. However, the correlation r_{corr} decreases with noise (middle panel of **Figure 2**) and the regression dilution bias still occurs with the linear-regression slopes decreasing with noise amplitude (bottom panel of **Figure 2**) interpreted as lessened Earth reaction for an increased driver strength.

As a preview of future work, adding noise to the solar-wind variables in real data [i.e., OMNI2, King and Papitashvili (2005)] indeed changes the functional form of the best-fit solar-wind driver. Fits of the form $v_{sw}^a B_{sw}^b \sin^c(\theta_{clock}/2)$ to various time-lagged geomagnetic indices (AE, AL, AU, Kp, Hp60, PCI) find that adding noise to any one of the three solar-wind variables changes the best-fit values of all three exponents a , b , and c . Depending on the geomagnetic index that is being fit, the best-fit values of a , b , or c can either decrease with added noise or increase with added noise. In agreement with the triangle points in the top panel of **Figure 2**, adding noise only to the geomagnetic index does not change the best-fit values of a , b , or c in a real data set. Real solar-wind data will be explored in a future report.

SUMMARY

The functional form obtained for the best fit solar-wind driver d depends on (at least) two things. It is a function of how the driving works. It is also a function of noise in the measurements. If our goal is to use real solar-wind/magnetosphere data to uncover or to confirm the formula that tells us the physics of the driving, we have trouble because of there always being noise in the data. One

source of error in the solar-wind and magnetosphere data is the fact that the solar wind that hits an upstream monitor is not the same solar wind that hits the earth: this error has been expounded upon (Borovsky, 2018; Borovsky, 2020b; Walsh et al., 2019; Burkholder et al., 2020). Another source of error is that geomagnetic indices are only indirect measures of the reaction of the earth to the solar wind. A future research effort might involve 1) obtaining a best-fit driver formula from the real data, 2) assessing the amplitude and properties of the noise in the real data, and 3) attempting to correct the formula for the effects of the noise.

DATA AVAILABILITY STATEMENT

The original contributions presented in the study are included in the article/Supplementary Material, further inquiries can be directed to the corresponding author.

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AUTHOR CONTRIBUTIONS

The author confirms being the sole contributor of this work and has approved it for publication.

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