



# Magnetic Field Reconstruction for a Realistic Multi-Point, Multi-Scale Spacecraft Observatory

T. Broeren<sup>1\*</sup>, K. G. Klein<sup>2</sup>, J. M. TenBarge<sup>3</sup>, Ivan Dors<sup>4</sup>, O. W. Roberts<sup>4</sup> and D. Verscharen<sup>5,6</sup>

<sup>1</sup>Department of Applied Mathematics, University of Arizona, Tucson, AZ, United States, <sup>2</sup>Lunar and Planetary Laboratory, University of Arizona, Tucson, AZ, United States, <sup>3</sup>Department of Astrophysical Sciences, Peyton Hall, Princeton University, Princeton, NJ, United States, <sup>4</sup>Space Research Institute, Austrian Academy of Sciences, Graz, Austria, <sup>5</sup>Mullard Space Science Laboratory, University College London, Surrey, United Kingdom, <sup>6</sup>Space Science Center, University of New Hampshire, Durham, NH, United States

## OPEN ACCESS

### Edited by:

Georgios Balasis,  
National Observatory of Athens,  
Greece

### Reviewed by:

Bogdan Hnat,  
University of Warwick,  
United Kingdom  
Arnaud Masson,  
European Space Astronomy Centre  
(ESAC), Spain  
Uwe Motschmann,  
Technische Universitat Braunschweig,  
Germany

### \*Correspondence:

T. Broeren  
broeren@email.arizona.edu

### Specialty section:

This article was submitted to  
Space Physics,  
a section of the journal  
Frontiers in Astronomy and Space  
Sciences

**Received:** 18 June 2021

**Accepted:** 13 August 2021

**Published:** 13 September 2021

### Citation:

Broeren T, Klein KG, TenBarge J M,  
Dors I, Roberts OW and Verscharen D  
(2021) Magnetic Field Reconstruction  
for a Realistic Multi-Point, Multi-Scale  
Spacecraft Observatory.  
Front. Astron. Space Sci. 8:727076.  
doi: 10.3389/fspas.2021.727076

Future *in situ* space plasma investigations will likely involve spatially distributed observatories comprised of multiple spacecraft, beyond the four and five spacecraft configurations currently in operation. Inferring the magnetic field structure across the observatory, and not simply at the observation points, is a necessary step towards characterizing fundamental plasma processes using these unique multi-point, multi-scale data sets. We propose improvements upon the classic first-order reconstruction method, as well as a second-order method, utilizing magnetometer measurements from a realistic nine-spacecraft observatory. The improved first-order method, which averages over select ensembles of four spacecraft, reconstructs the magnetic field associated with simple current sheets and numerical simulations of turbulence accurately over larger volumes compared to second-order methods or first-order methods using a single regular tetrahedron. Using this averaging method on data sets with fewer than nine measurement points, the volume of accurate reconstruction compared to a known magnetic vector field improves approximately linearly with the number of measurement points.

**Keywords:** plasma physics, magnetic fields, spacecraft, vector field reconstruction, space physics, curlometer, space mission analysis, multi-spacecraft analysis

## 1 INTRODUCTION

Plasmas, which are ubiquitous throughout the Universe, are readily available for study in the natural laboratory of space. Electromagnetic fields play a fundamental role in the transport, heating, and acceleration of charged particles that compose plasmas. In order to characterize fundamental processes governing heliospheric plasmas, the space plasma community has utilized *in-situ* spacecraft measurements of electromagnetic fields and charged particles. These *in-situ* measurements include the characterization of the vector magnetic field  $\mathbf{B}$  at a spacecraft via magnetometers; see §2.4 of Verscharen et al. (2019).

Knowledge of  $\mathbf{B}$  from a single magnetometer is limited; single-point measurements can not construct the full three-dimensional structure characteristic of processes such as magnetic reconnection and plasma turbulence. To avoid this shortcoming, ESA's CLUSTER (Escoubert et al., 2001), NASA's THEMIS (Angelopoulos, 2008) and MMS (Burch et al., 2016) missions have employed four- and five-spacecraft configurations, where each spacecraft is equipped with an

instrument suite that includes a magnetometer. These missions study the boundaries of the Earth's magnetosphere, including how magnetic reconnection transfers magnetic energy into kinetic energy of plasma particles.

Analysis techniques have been created for multi-spacecraft missions, such as CLUSTER, which search for specific types of plasma waves (Constantinescu et al., 2006) and which analyze current sheet structure (Narita et al., 2013) for a four spacecraft configuration. Knowledge of the direction of wave propagation allows us to use multi-spacecraft filtering (Pinçon and Motschmann, 1998) to determine the general polarisation properties of any multi-point measurement of a wave field in space plasmas. Measurements from exactly four spacecraft (e.g., a tetrahedron of spacecraft) can be used to estimate current density via the Curlometer technique (Robert et al., 1998).

The Cluster and MMS missions have also utilized the Curlometer technique to interpolate the value of the magnetic field over a region near the tetrahedron's barycenter, regardless of the field's geometry. However, these interpolations are limited to measuring fluctuations on a scale on the same order as that of their inter-spacecraft distances (e.g., Robert et al., 1998; Forsyth et al., 2011). To study multi-scale processes, such as plasma turbulence, with structures on characteristic length scales that cover many orders of magnitude, we must employ measurements from more than four spacecraft. Therefore, we develop a method which extends the magnetic field reconstruction technique Curlometer to configurations of more than four spacecraft.

Many such multi-spacecraft missions have been proposed, e.g., Cross-scale (Schwartz et al., 2009), AME (Dai et al., 2020) and HelioSwarm (Klein et al., 2019), but in order to optimize such missions, it is urgent to robustly quantify the impact of particular spacecraft configurations on multi-point analysis methods, capturing the effects of the physical scales spanned by the spacecraft in the observatory and the geometry of the polyhedra that can be drawn from the constituent spacecraft. Such quantification will help demonstrate that a proposed mission will be able to usefully analyze a large number of magnetometer measurements made in the pristine solar wind, magnetosphere, and magnetosheath. It will also assist in the optimization of spacecraft configurations and quantification of errors derived from multi-point, multi-scale measurements. In this paper, we focus on the fidelity of the reproduction of the magnetic field using a sparsely sampled set of measurements whose spatial configuration is based upon realistic configurations of the proposed nine-spacecraft HelioSwarm observatory, described for instance by Plice et al. (2020).

The reconstruction method is described in §2, the results are applied to two magnetic field models, including a numerical simulation of turbulence, in §3, with a concluding discussion in §4.

## 2 METHODOLOGY

### 2.1 Geometrical Definitions

Given  $N$  spacecraft, we identify  $C(N, k)$  polyhedra with  $k$  vertices. As spatial divergence analysis methods (e.g., Dunlop et al., 1988;

Paschmann and Daly, 1998, 2008) require at least four vertices to resolve three-dimensional structure, we only consider polyhedra with at least four vertices, known as tetrahedra. For  $N = 9$ , there are 126 (i.e., 9 choose 4) tetrahedra, 126 polyhedra with five vertices, 84 with six vertices, 36 with seven vertices, 9 with 8 vertices, and 1 with 9 vertices, for a total of 382 polyhedra with at least four vertices.

Each polyhedron is characterized in terms of its size and shape. Because measurements from all  $d$  spacecraft are weighted equally, we define the *barycenter* of the  $q$ th polyhedron with set  $\mathbf{D}$  of  $d$  vertices drawn from the  $N \geq d$  spacecraft positions  $\mathbf{x}_i$  as

$$\mathbf{x}_0^{q,d} = \frac{1}{d} \sum_{i \in \mathbf{D}} \mathbf{x}_i. \quad (1)$$

Given the barycenter, we then define the volumetric tensor of the  $q$ th polyhedra with set  $\mathbf{D}$  of  $d$  vertices as

$$\mathbf{R}_{jk}^{q,d} = \frac{1}{d} \sum_{i \in \mathbf{D}} (x_{ij} - x_{0j}^{q,d})(x_{ik} - x_{0k}^{q,d}). \quad (2)$$

Here  $x_{ij}$  represents the  $j$ th  $\in \{x, y, z\}$  component of the position vector for the  $i$ th spacecraft. The eigenvectors of the tensor  $\mathbf{R}^{q,d}$  represent the three semi-axes of the polyhedra and are associated with the eigenvalues  $a^{q,d} = \sqrt{R_1^{q,d}}$  (major axis),  $b^{q,d} = \sqrt{R_2^{q,d}}$  (middle axis), and  $c^{q,d} = \sqrt{R_3^{q,d}}$  (minor axis), where  $a \geq b \geq c$  (a more detailed analysis of the eigenvalues can be found in Chapt 12 of Paschmann and Daly, 1998).

To provide a useful geometric interpretation of these shapes, we define a *characteristic size*  $L$ , as well as an *elongation*  $E$  and a *planarity*  $P$  (see chapter 16.3 of Paschmann and Daly, 1998)<sup>1</sup>:

$$\begin{aligned} L &= 2a \\ E &= 1 - b/a \\ P &= 1 - c/b. \end{aligned} \quad (3)$$

## 2.2 Reconstruction Techniques

### 2.2.1 First-Order Method

In a first-order Taylor series expansion, we use the values of the magnetic field,  $\mathbf{B}$ , measured at four spacecraft positions  $\mathbf{x}_i$  to estimate the value of  $\mathbf{B}$  (and its corresponding directional derivatives) at any other point in space,  $\xi$  (Fu et al., 2015, 2020). The Taylor expansion is:

$$\hat{B}_m^i = B_m + \sum_{k \in \{x,y,z\}} \partial_k B_m r_k^i \forall i \in \{1, 2, 3, 4\}, m \in \{x, y, z\}. \quad (4)$$

In this equation  $\hat{B}_m^i$  is the measured  $m$ th component of  $\mathbf{B}$  at the  $i$ th spacecraft,  $B_m$  is the computed  $m$ th component of  $\mathbf{B}$  at  $\xi$ ,  $\partial_k B_m$  is the computed derivative of the  $m$ th component of  $\mathbf{B}$  with respect to the  $k$ th direction at  $\xi$ , and  $r_k^i$  is the relative position of

<sup>1</sup>Note that there is some discrepancy in the community about how the elongation should be defined as  $E = 1 - b/a$  or  $E = 1 - 1/(b/a)^2$ , with a similarly subjective choice for planarity. Both definitions span the same range, and we have opted for the former definition.

spacecraft  $i$  with respect to  $\xi$ . In other words, if  $x_{ik}$  is the  $k$ th component of spacecraft  $i$ 's location, then  $r_k^i := x_{ik} - \xi_k$ .

This is a system of 12 equations with 12 unknowns, where the 12 equations represent the  $x$ ,  $y$ , and  $z$  components of  $\mathbf{B}$  for each of the four spacecraft. The 12 unknowns are the  $x$ ,  $y$ , and  $z$  components of  $\mathbf{B}$  at  $\xi$  and the nine terms in the Jacobian of  $\mathbf{B}$  at  $\xi$ .

This system can be reformatted into linear ( $A\mathbf{x} = \mathbf{b}$ ) form and solved with a common linear system solver. This 12-dimensional linear system (shown in full detail in **Supplementary Appendix SA**, Eq. A8) comprises the first-order reconstruction method.

This magnetic field reconstruction method is related to the Curlometer method (Dunlop et al., 1988; Robert et al., 1998), which utilizes Ampère's law to calculate the current density  $\mathbf{J}$  as the curl of  $\mathbf{B}$ . The Curlometer solves the same set of equations, but uses the partial derivatives to estimate the current density at the center of each tetrahedron. The Curlometer method has been widely applied to four-spacecraft magnetic field measurements made for instance by Cluster and MMS (c.f. Chapter 16.2 of Paschmann and Daly, 1998). Future missions, such as the proposed HelioSwarm Observatory (Klein et al., 2019), will have more than four spacecraft. Therefore, for every reconstructed point,  $\xi$ , we can apply this reconstruction method for each of the  $C(N, 4)$  tetrahedra and average the reconstructed values, yielding a statistically larger base of estimates and improving the accuracy of the reconstruction.

### 2.2.2 Second-Order Method

Because the proposed HelioSwarm Observatory has nine spacecraft, we can use measurements of  $\mathbf{B}$  from all nine spatial points simultaneously to apply a second-order reconstruction method. This method, also based on a Taylor series expansion, is more accurate for values located near the center of the expansion (i.e., near the barycenter of the nine-spacecraft constellation) than a single implementation of the first-order method. Following the work of Torbert et al. (2020), we write:

$$\hat{B}_m^i = B_m + \sum_{k \in \{x,y,z\}} \partial_k B_m r_k^i + \frac{1}{2} \sum_{j,k \in \{x,y,z\}} \partial_j \partial_k B_m r_k^i r_j^i \quad \forall i \in \{1, \dots, 9\}, m \in \{x, y, z\}. \tag{5}$$

These terms are the same as in the first-order method, with the addition of  $\partial_j \partial_k B_m$ , the second derivative of the  $m$ th component of  $\mathbf{B}$  with respect to the  $k$ th and  $j$ th directions, at  $\xi$ .

This is a system of 31 equations with 30 unknowns. 27 of these equations are associated with the  $x$ ,  $y$ , and  $z$  components of  $\mathbf{B}$  from the nine spacecraft. There are four additional constraints, imposed by the magnetic field having zero divergence, as well as the divergence of the magnetic field having zero gradient. The 30 unknowns are the  $x$ ,  $y$ , and  $z$  components  $\mathbf{B}$  at  $\xi$ , the nine terms in the Jacobian of  $\mathbf{B}$  at  $\xi$ , and the 18 terms in the Hessian of  $\mathbf{B}$  at  $\xi$  (excluding the 9 redundant terms).

This system can be reformatted into linear ( $A\mathbf{x} = \mathbf{b}$ ) form where  $A$  is a  $31 \times 30$  matrix (shown in full detail in **Supplementary Appendix SB**, Eq. B16). This system is over-determined, therefore in general, an exact solution does not exist. However, we can find an approximate solution via the method of

ordinary least squares. This method finds the solution to the problem  $A\mathbf{x} = \mathbf{b}$  which minimizes the two-norm of the error, i.e.

$$\mathbf{x} = \underset{\mathbf{x}}{\operatorname{argmin}} \|A\mathbf{x} - \mathbf{b}\|_2. \tag{6}$$

This second-order reconstruction method is referred to as  $M_2$  throughout this paper.

### 2.2.3 Quantifying Error

We define the error at any point in space,  $\xi$ , as:

$$\theta = 100 \frac{\|\mathbf{B}_{\text{calc}}(\xi) - \mathbf{B}_{\text{true}}(\xi)\|_2}{\|\mathbf{B}_{\text{true}}(\xi)\|_2} \tag{7}$$

where  $\mathbf{B}_{\text{true}}(\xi)$  is the magnetic field vector at point  $\xi$  and  $\mathbf{B}_{\text{calc}}(\xi)$  is the reconstructed magnetic field vector at point  $\xi$ .

Given that we can determine the value of this error at all points in a simulation or for a given analytic field, we also define  $\epsilon(\theta)$  as the proportion of the volume that is reconstructed with less than  $\theta\%$  error. For a sufficiently dense grid of  $N_x \times N_y \times N_z$  uniformly-spaced points,  $\epsilon(\theta)$  can be estimated as

$$\epsilon(\theta) = \frac{\#\text{points with } \leq \theta\% \text{ error}}{N_x N_y N_z}. \tag{8}$$

To define the physical volume in which a given magnetic field reconstruction is accurate,  $\epsilon(\theta)$  is translated into a dimensional quantity by multiplying it by the total volume covered by the  $N_x \times N_y \times N_z$  grid.

### 2.2.4 Error Minimization Techniques

As the first-order method (§2.2.1) only requires a single tetrahedron of spacecraft to estimate  $\mathbf{B}$ , in this paper we will test four selection methods for using a subset of the 126 tetrahedra to improve the reconstruction. These methods combine the statistically large set of tetrahedra with our knowledge of the spacecraft positions relative to  $\xi$  and the geometry of all 126 tetrahedra.

For the first method,  $M_{1,1}$ , at each point in space we reconstruct the magnetic field using all 126 tetrahedra to produce 126 estimates for  $\mathbf{B}(\xi)$ . We then average over these  $\mathbf{B}(\xi)$  values component-wise to estimate  $\mathbf{B}(\xi)$ .

For the second method,  $M_{1,2}$ , we perform the same averaging as method one, but only include tetrahedra whose barycenter are within a characteristic distance of  $\xi$ . i.e., for each reconstructed point  $\xi$ , only include tetrahedra  $j$  in the average which satisfy

$$\|(r_0)_j - \xi\|_2 < L_j, \tag{9}$$

where  $L_j$  is the characteristic size and  $(r_0)_j$  is the barycenter of the  $j$ th tetrahedron.

For the third method,  $M_{1,3}$ , we perform the same selection as method two, but with the added restriction that the shape of tetrahedron  $j$  must be quasi-regular. In terms of the geometric quantities of the spacecraft configuration (defined in **Eq. 3**), this translates to elongation  $E$  and planarity  $P$  being sufficiently small. Because  $E$  and  $P$  are symmetric with respect to orientation, we will define a composite geometric parameter  $\chi_j$

$$\chi_j = \sqrt{E_j^2 + P_j^2}. \quad (10)$$

Small  $\chi_j$  implies that both the elongation and planarity of tetrahedron  $j$  are small. For method  $M_{1,3}$ , we restrict our averaging to only include tetrahedra where

$$\| (r_0)_j - \xi \|_2 < L_j \quad \chi_j \leq 1 \quad (11)$$

For the fourth method,  $M_{1,4}$ , we perform the same selection of tetrahedra as method three, but require the tetrahedra included in the averaging to be more regular. For method  $M_{1,4}$ , our shape and position requirement is

$$\| (r_0)_j - \xi \|_2 < L_j \quad \chi_j \leq 0.6 \quad (12)$$

The value of 0.6 was selected because page 408 of Paschmann and Daly (1998) shows it be a threshold value for elongation and planarity which separates the well performing ‘pseudo-sphere type’ and ‘potato type’ spacecraft configurations from the poorer performing ‘knife blade type’, ‘cigar type’, and ‘pancake type’ configurations.

The first-order methods  $M_{1,1}$ ,  $M_{1,2}$ ,  $M_{1,3}$ ,  $M_{1,4}$  will be compared to the second-order method  $M_2$  as well as the first-order method applied to a single regular (i.e.,  $\chi = 0$ ) tetrahedron of spacecraft. This single regular tetrahedron will have the same characteristic scale as the nine-spacecraft configuration it is compared to.

### 2.3 Models

To validate and quantify the errors of our reconstruction, we implement our reconstruction methods on two magnetic field models, a simple current sheet and a numerical simulation of turbulence.

#### 2.3.1 Simple Current Sheet

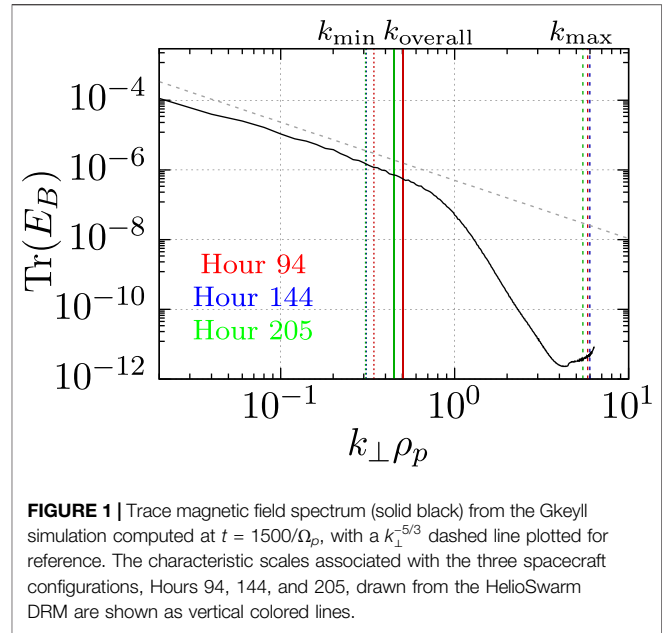
For our first model, we define a magnetic field where  $\mathbf{B}$  is analytically defined at all spatial points. This field, which represents a simple current sheet, can be described in cylindrical coordinates as

$$\mathbf{B}(r) = \mu_0 J_0 \sigma [\sigma/r - e^{-r\sigma} (1 + \sigma/r)] \hat{\theta}. \quad (13)$$

The variable  $\sigma$  represents the current sheet characteristic width and  $J_0$  represents the magnitude of the current at its center.

#### 2.3.2 Turbulence Simulation

Physically realistic fields, such as those generated by turbulence in the solar wind, are significantly more complex than the simple current sheet model of Eq. 13. We therefore test our reconstruction techniques on magnetic fields drawn from numerical simulations of turbulence. In particular, we utilize the magnetic fields from a fully developed turbulence simulation performed with the five moment, multi-fluid solver within the Gkeyll simulation framework (Hakim et al., 2006;



**FIGURE 1** | Trace magnetic field spectrum (solid black) from the Gkeyll simulation computed at  $t = 1500/\Omega_p$ , with a  $k_{\perp}^{-5/3}$  dashed line plotted for reference. The characteristic scales associated with the three spacecraft configurations, Hours 94, 144, and 205, drawn from the HelioSwarm DRM are shown as vertical colored lines.

Wang et al., 2015, 2020). This turbulence simulation is designed to represent plasma behavior in the pristine solar wind at 1AU.

We use the five moment ( $n_s, \mathbf{u}_s, p_s$ ), two fluid ( $s = p, e$ ) plasma model to evolve a proton-electron plasma. We note that the five moment, two fluid model formally reduces to Hall MHD in the limit  $m_e \rightarrow 0$  and  $\epsilon_0 \rightarrow 0$  (Srinivasan and Shumlak, 2011), where  $\epsilon_0$  is the vacuum permittivity. We use a reduced (proton to electron) mass ratio of  $m_p/m_e = 100$ , a temperature ratio of  $T_p/T_e = 1$ , Alfvén velocity of  $v_A/c = B/\sqrt{\mu_0 n_p m_p c^2} = 0.02$ , plasma beta (ratio of plasma thermal pressure to magnetic pressure) of  $\beta_p = 2\mu_0 n_p T_p/B^2 = 1$ , and adiabatic index  $\gamma = 5/3$ . We employ an elongated domain  $L_x = L_y = 0.2L_z = 100\pi\rho_p$  with resolution  $n_x = n_y = n_z = 448$ . Lengths are normalized to the proton gyroradius  $\rho_p = v_{tp}/\Omega_p$ , the ratio of the proton thermal speed  $v_{tp} = \sqrt{2T_p/m_p}$  and the proton cyclotron frequency  $\Omega_p = q_p B/m_p$ . We choose a uniform background density and magnetic field,  $\mathbf{B}_0 = B_0 \hat{z}$ , and initialize the simulation with the three dimensional extension of the Orszag-Tang vortex (Orszag and Tang, 1979) described in Li et al. (2016).

$$\begin{aligned} \frac{\mathbf{z}_1^+}{v_A} &= -\frac{2z_0}{v_A} \sin(k_{\perp}y - k_z z) \hat{\mathbf{x}}, \quad \frac{z_1^-}{v_A} = 0 \\ \frac{\mathbf{z}_2^+}{v_A} &= \frac{z_0}{v_A} \sin(k_{\perp}x \mp k_z z) \hat{\mathbf{y}} \\ \frac{\mathbf{z}_3^+}{v_A} &= \pm \frac{z_0}{v_A} \sin(2k_{\perp}x \mp k_z z) \hat{\mathbf{y}}, \end{aligned} \quad (14)$$

where  $\mathbf{z}^{\pm} = \delta\mathbf{u} \pm \delta\mathbf{B}/\sqrt{\mu_0\rho_0}$  are the Elsasser variables (Elsasser, 1950),  $k_{x,y} = 2\pi/L_{x,y}$ , and  $k_z = 2\pi/L_z$ . The initial amplitude,  $z_0 = 0.2$ , is chosen to satisfy the critical balance condition,  $k_x z_0/k_z v_A = 1$  (Goldreich and Sridhar, 1995).

The simulation is run for one Alfvén crossing time,  $t_A = 1500/\Omega_p$ , at which point the turbulence has fully developed and reached

**TABLE 1 |** Characteristic geometric parameters for the three nine-spacecraft configurations under consideration and the minimum/maximum characteristic geometric parameters created from choosing any four of the nine spacecraft of each configuration.

Hour		L (km)	E	P
94	Overall	1,245	0.48	0.60
	Min	108	0.14	0.10
	Max	1834	0.93	0.99
144	Overall	1,395	0.42	0.70
	Min	108	0.06	0.23
	Max	2030	0.95	0.99
205	Overall	1,401	0.45	0.75
	Min	115	0.32	0.26
	Max	2045	0.97	0.99

a steady state. In **Figure 1**, we plot the trace magnetic energy spectrum as a function of  $k_{\perp} = \sqrt{k_x^2 + k_y^2}$ , with a  $k_{\perp}^{-5/3}$  dashed line plotted for reference. The steep roll-over in the spectrum at  $k_{\perp} \rho_p \approx 1$  is due to numerical diffusion from the finite volume scheme employed by Gkeyll.

To compare the simulation to the selected spacecraft configurations with separations in physical units, we note that the proton gyroradius can be written as

$$\rho_p = \sqrt{\frac{m_p}{m_e}} \sqrt{\beta_p} \frac{c}{\omega_{pe}} \tag{15}$$

With the constants in the turbulence simulation of  $m_p/m_e = 100$ ,  $\beta_p = 1$ ,  $\omega_{pe} = \frac{c}{d_e} = 5.64 \times 10^4 \sqrt{n}$ , we set  $n_e = 0.2829 \text{ cm}^{-3}$ , so that  $\rho_p = 100 \text{ km}$ .

We extract from this simulation a 3-dimensional grid of values representing the plasma’s physical parameters at different points in space. From this grid, we use trilinear interpolation to estimate the value of **B** at any point in the simulation volume.

### 2.3.3 Spacecraft Configurations

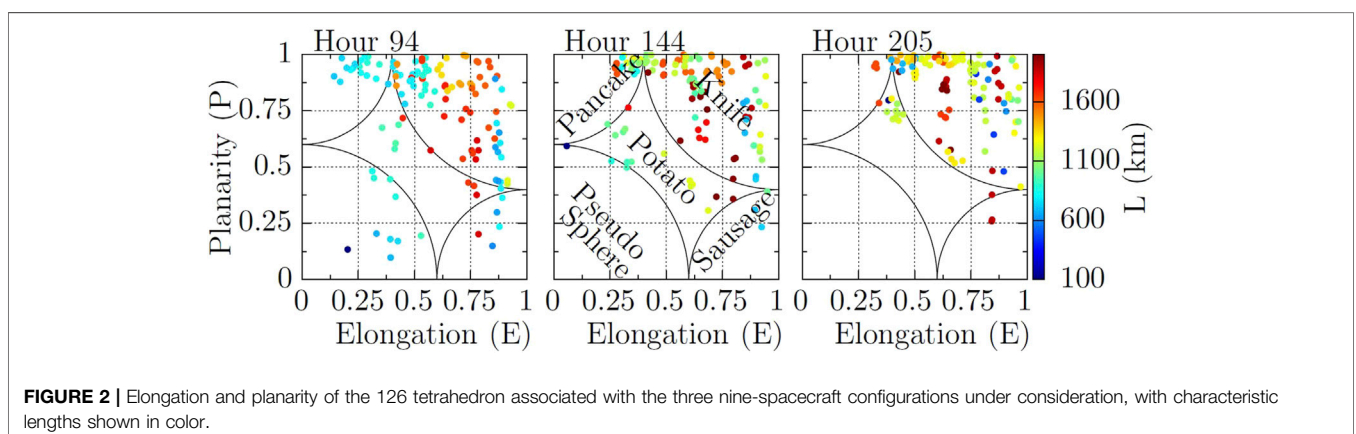
To illustrate our reconstruction methods for realistic spacecraft configurations, we study these methods using three different nine-spacecraft configurations. The spacecraft configurations are selected from the phase A

design reference mission (DRM) of the proposed HelioSwarm Observatory concept, corresponding to hours 94, 144, and 205 of the science phase. These hours are selected because they represent a selection of spacecraft tetrahedra that have significantly different distributions of their elongation, planarity, and length. In **Table 1** we note the geometric characteristics of the overall nine-vertex polyhedra for each of the three configurations. We also calculate the size, elongation, and planarity of all 126 tetrahedron in each configuration and display them in **Figure 2**, noting the minimum and maximum values of these three parameters for each configuration in **Table 1**. The wavelengths associated with the overall, minimum, and maximum scales,  $k\rho_p = 2\pi\rho_p/L$ , are overlaid on **Figure 1**, using a fiducial value of  $\rho_p = 100 \text{ km}$ .

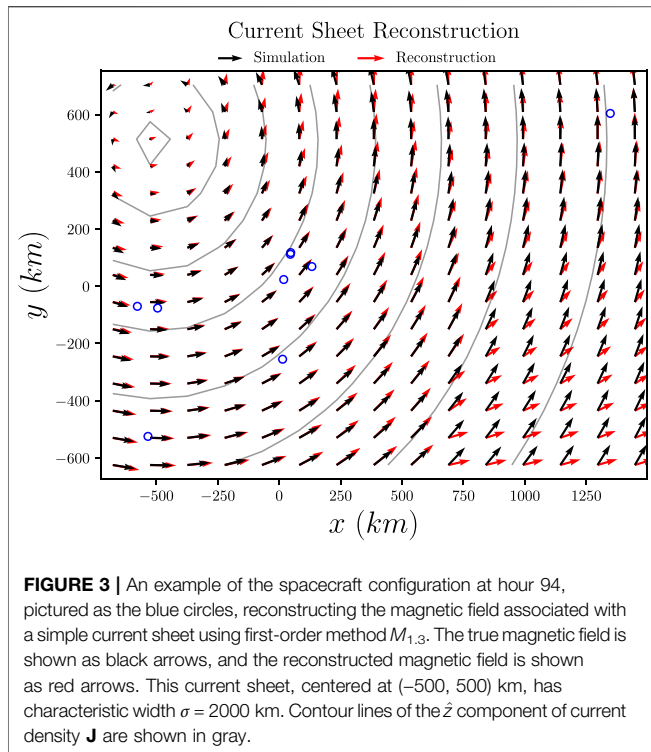
## 3 APPLICATION OF RECONSTRUCTION

To find the expected error at all points in space near a particular spacecraft configuration, we take a Monte Carlo approach and place the barycenter of each nine-spacecraft configuration into a known magnetic field at random locations. We then reconstruct the magnetic field on a grid of points centered at the barycenter of the nine-spacecraft configuration using the first- and second-order reconstruction methods. The location of each point in the reconstructed grid is constant with respect to the spacecraft configuration. Therefore, we find the average of the errors,  $\theta$ , at all reconstructed grid points for all elements of the Monte Carlo ensemble, allowing the calculation of the expected value of error at each point on the grid.

Additionally, we compare the divergence found on a grid of points sampled from the baseline current sheet and turbulence simulation magnetic fields with that of the same points sampled from the fields reconstructed using our first-order reconstruction methods. This comparison yields divergence values of similar magnitude in the baseline and reconstructed fields, which indicates that our reconstruction methods do not introduce nonphysical values of divergence.



**FIGURE 2 |** Elongation and planarity of the 126 tetrahedron associated with the three nine-spacecraft configurations under consideration, with characteristic lengths shown in color.

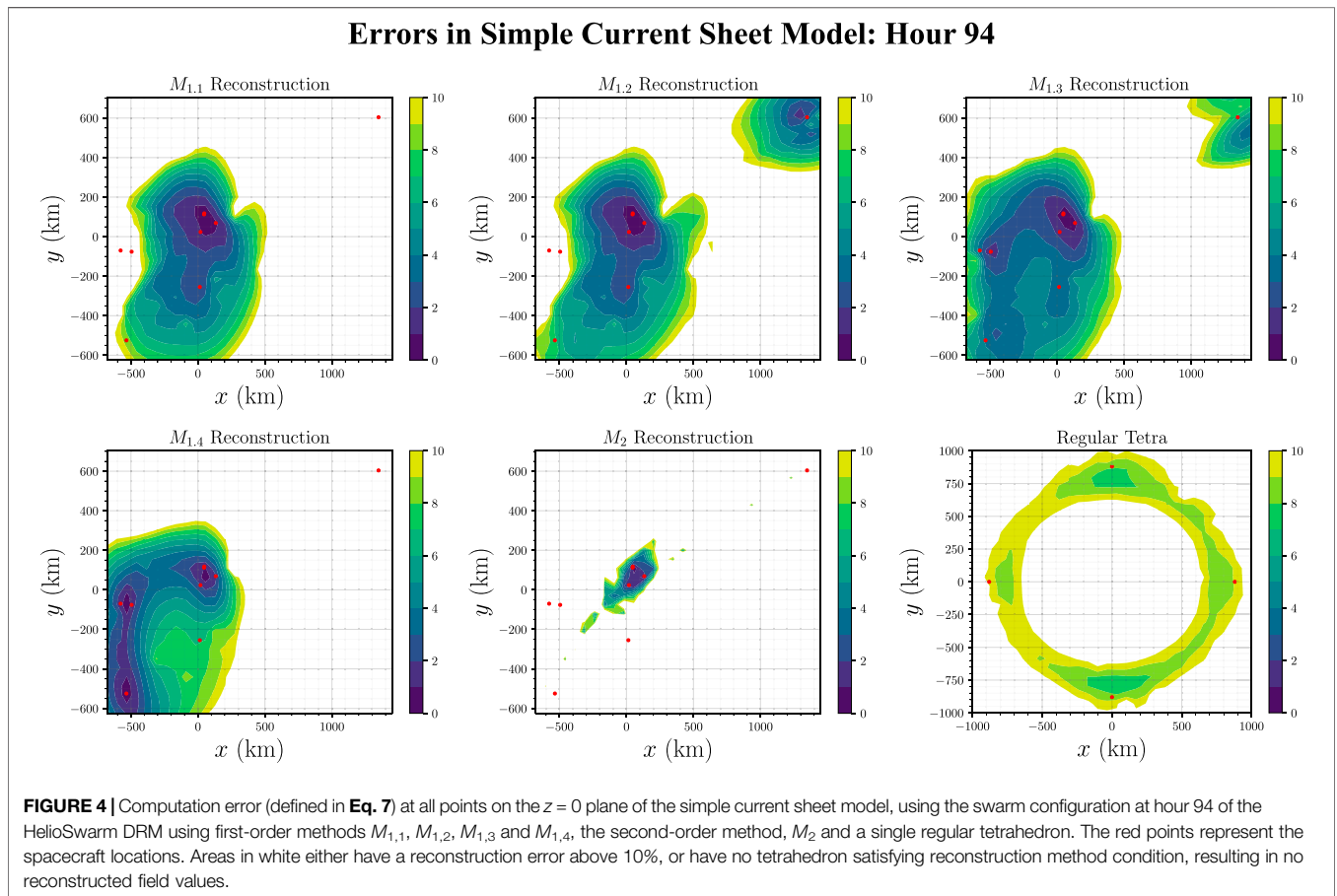


### 3.1 Current Sheet

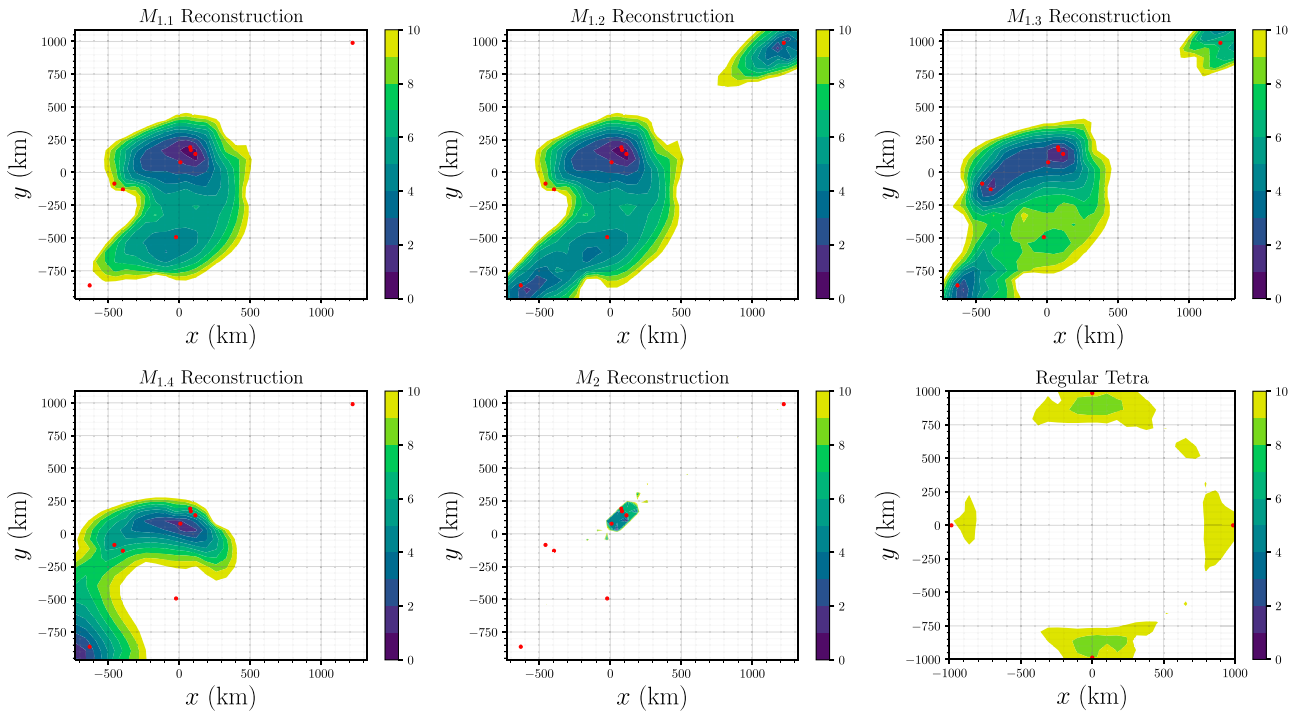
We present an example magnetic field reconstruction of the simple current sheet model (§2.3.1) in **Figure 3**. Here, we use the first-order method  $M_{1,3}$  to reconstruct the magnetic field in the  $z = 0$  plane for the simple current sheet, **Eq. 13** with  $\sigma = 2000$  km using the hour 94 spacecraft configuration. There is little difference between the reconstructed and original fields near the center of the spacecraft configuration, and the difference in vectors increases with distance from the spacecraft configuration’s center.

We perform 200 Monte Carlo iterations of reconstruction using each method, observing that 200 was sufficient to point-wise converge in error. The characteristic width of the current sheet is chosen as a uniform random variable  $\sigma \sim U [500, 5000]$  km, while the barycenter of the nine-spacecraft configuration is selected as a 3D uniform random variable  $r_0 \sim U [-1000, 1000]^3$  km. We reconstruct a  $30 \times 30 \times 30$  grid of points  $\xi$  that extends 100 km past the furthest spacecraft in all directions.

The errors computed for the reconstruction of the simple current sheet are displayed in **Figures 4–6** for the hour 94, 144, and 205 configurations respectively. These figures illustrate the ensemble-averaged errors along a 2D plane orthogonal to the current intersecting a given nine-spacecraft configuration’s barycenter. The first four panels correspond to the four first-order reconstruction methods,  $M_{1,1}$ ,  $M_{1,2}$ ,  $M_{1,3}$  and  $M_{1,4}$ , the fifth

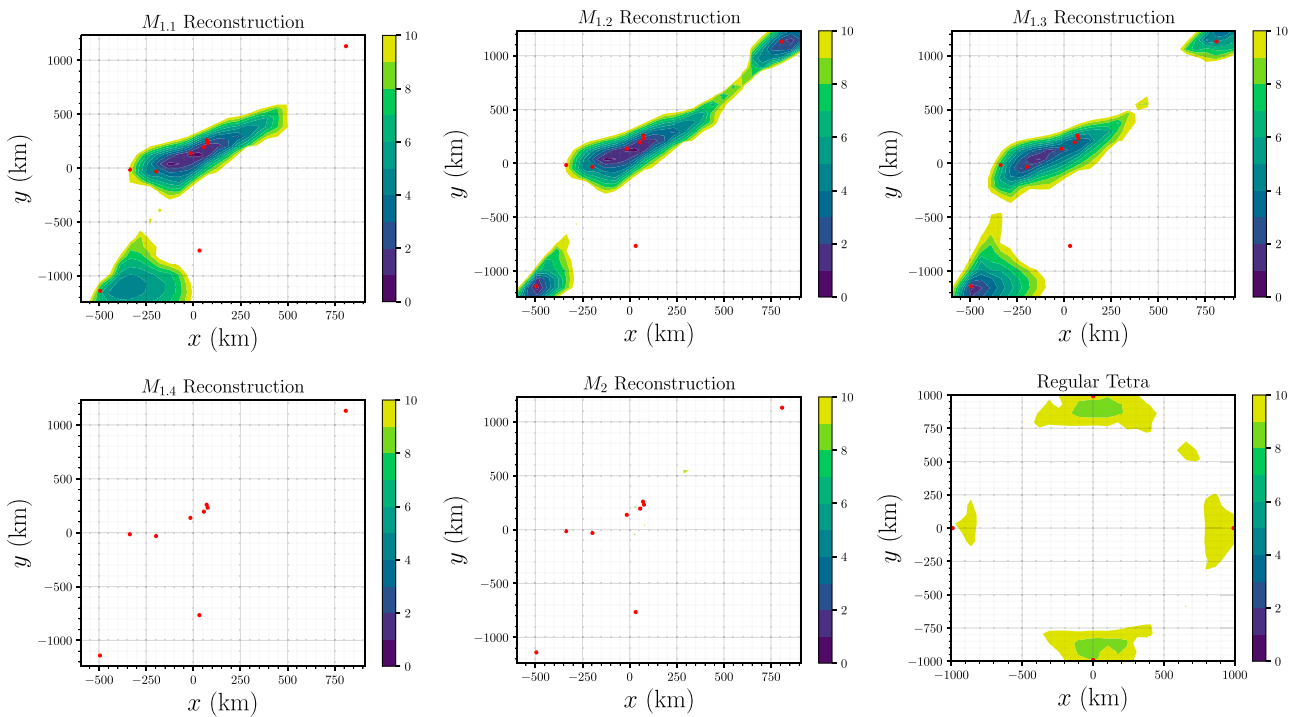


### Errors in Simple Current Sheet Model: Hour 144

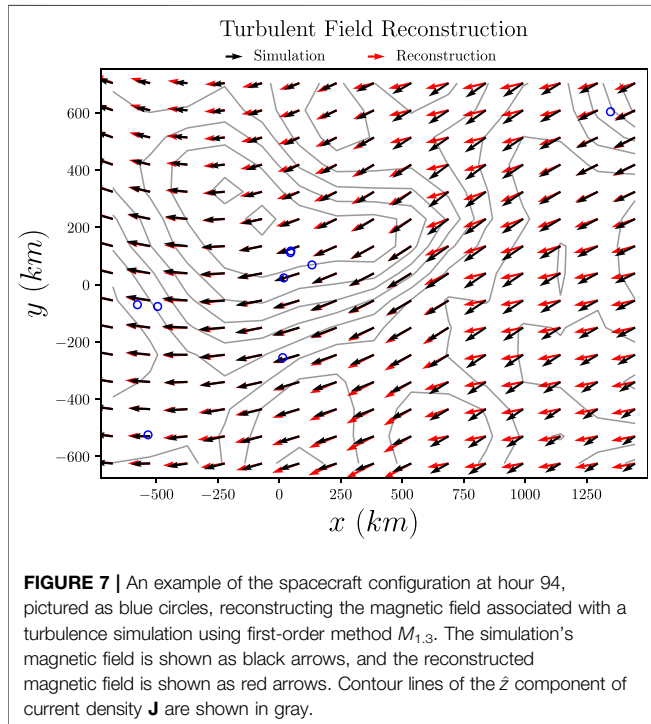


**FIGURE 5** | Computation error (defined in Eq. 7) at all points on the  $z = 0$  plane of the simple current sheet model, using the swarm configuration at hour 144 of the HeliSwarm DRM. The layout is identical to Figure 4.

### Errors in Simple Current Sheet Model: Hour 205

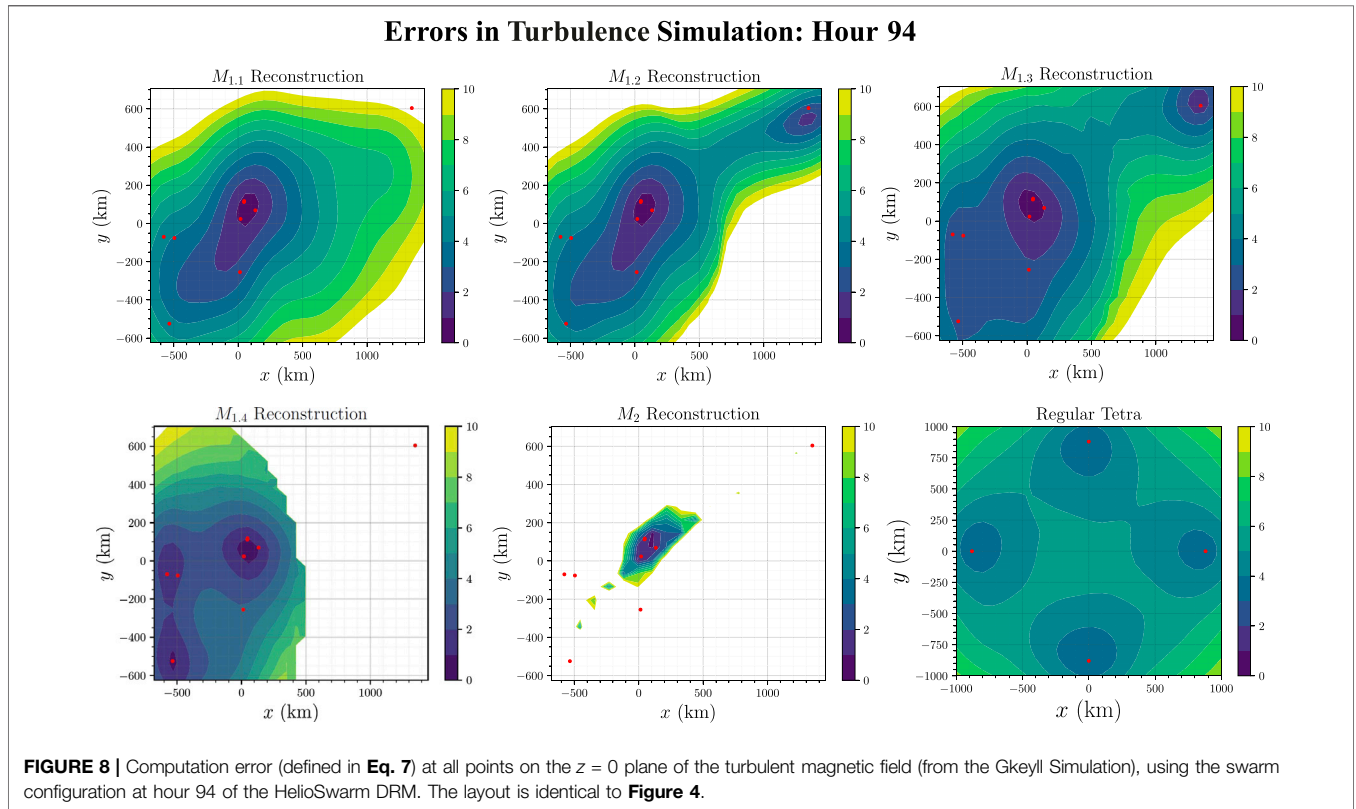


**FIGURE 6** | Computation error (defined in Eq. 7) at all points on the  $z = 0$  plane of the simple current sheet model, using the swarm configuration at hour 205 of the HeliSwarm DRM. The layout is identical to Figure 4.

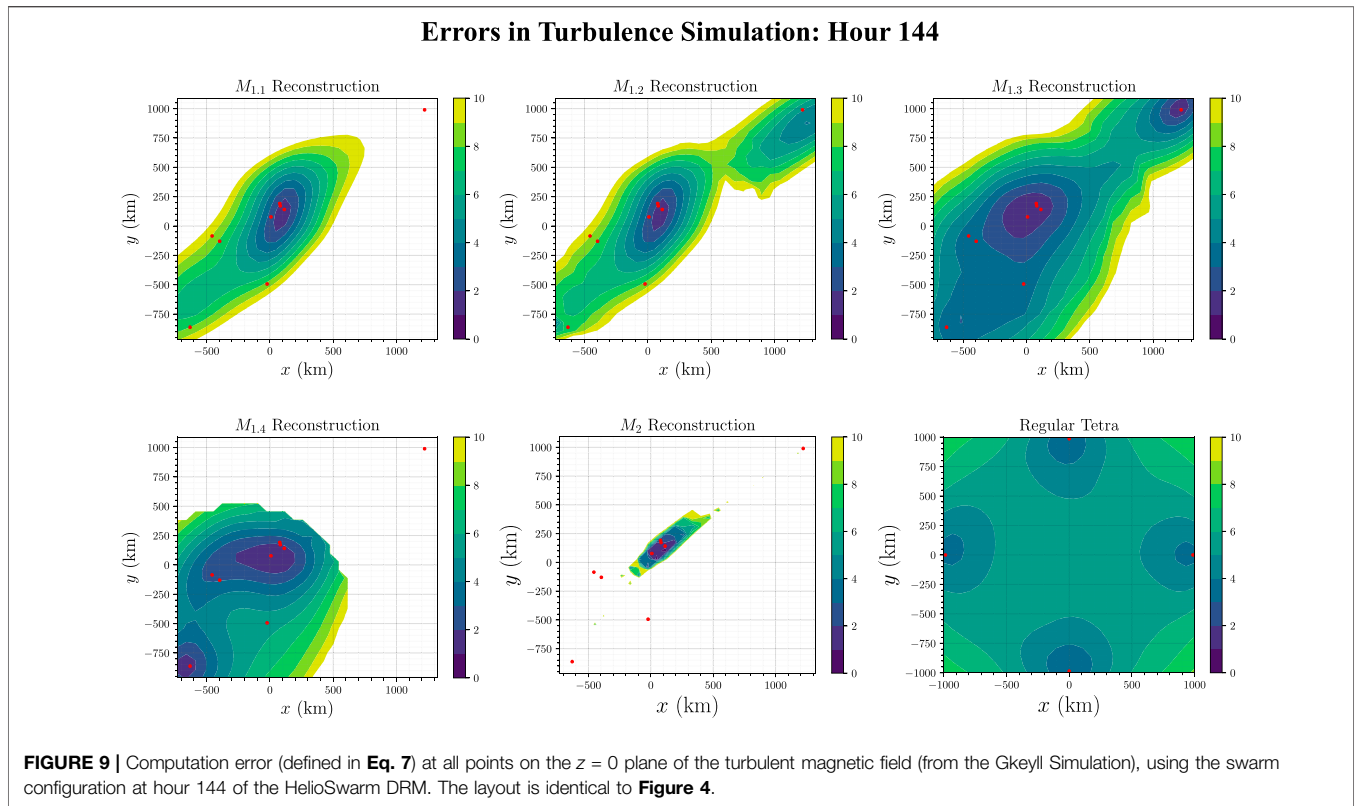


panel corresponds to the second-order method  $M_2$ , and the final panel corresponds to the reconstruction obtained from the standard first-order method applied to a single regular tetrahedron, with  $E = P = 0$ . This single tetrahedron has the same characteristic size  $L$  as the overall nine-spacecraft configuration, calculated as twice the major axis of the volumetric tensor, Eq. 2, evaluated using all nine points. With four spacecraft, we cannot reconstruct the magnetic field with the second-order method, nor can we select subsets of tetrahedra with advantageous geometric characteristics, so only the first-order reconstruction method from a single tetrahedron is used.

We see that near the barycenter of each of the nine-spacecraft configurations (located at the origin of **Figures 4–6**) the magnetic field can be reconstructed to within 1% accuracy. By comparing method  $M_{1,1}$  with methods  $M_{1,2}$  and  $M_{1,3}$  in these figures, we also conclude that leveraging knowledge of the tetrahedral shapes and positions expands the region of high-accuracy reconstruction. Unfortunately, overly restrictive conditions limit the number of tetrahedra available to average over, limiting the size of the reconstructed region. In fact, the bottom left panel of **Figure 6** is empty because none of the 126 tetrahedra in the hour 205 configuration satisfy the geometric requirement that  $\chi_j \leq 0.6$  demanded by  $M_{1,4}$ .







Additionally, the second-order reconstruction method  $M_2$  is accurate for only a small volume when compared with the first-order methods  $M_{1,1}$ ,  $M_{1,2}$ ,  $M_{1,3}$ , and  $M_{1,4}$ .

By comparing the bottom right panel to the other five in Figures 7–10, we see that the behavior of methods  $M_{1,1}$ ,  $M_{1,2}$ ,  $M_{1,3}$ ,  $M_{1,4}$ , and  $M_2$  is distinct to that of the reconstruction using a single regular tetrahedron. The single regular tetrahedron only accurately reconstructs the magnetic field of the current sheet near each of the four spacecraft. Due to the angular symmetry in the current sheet and the fact that none of the four spacecraft are positioned on the  $z = 0$  plane, the area of most accurate reconstruction appears to be a ring on the bottom right panel of Figures 4–6.

### 3.2 Turbulence Simulation

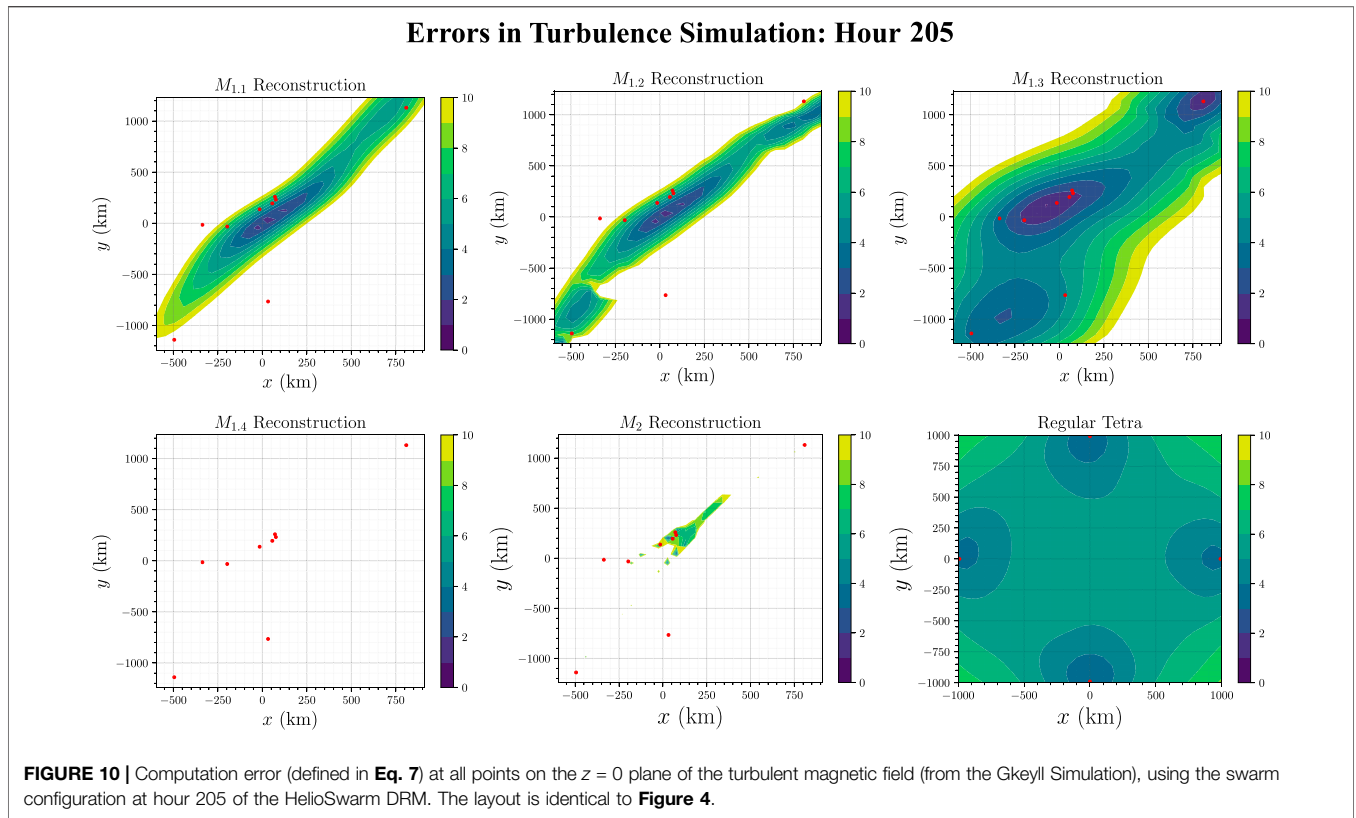
We present an example magnetic field reconstruction of the turbulence simulation (§2.3.2) in Figure 7. Here, we use the first-order method  $M_{1,3}$  to reconstruct the magnetic field in the  $z = 0$  plane in the turbulence simulation. Mirroring the behavior described in Figure 3, there is little difference between the reconstructed and original fields near the center of the spacecraft configuration.

We perform 50 Monte Carlo iterations of reconstruction using each method, observing that 50 was more than enough to pointwise converge in error. The barycenter is chosen as a uniform random variable so that all spacecraft remained in the  $31415 \times 31415 \times 157079$  km simulation cube. We then construct a  $30 \times 30 \times 30$  grid of points  $\xi$ . Each dimension of this grid is selected so

that the overall size of the grid extends 100 km past the furthest spacecraft in all directions.

The errors computed from the turbulence simulation reconstruction are displayed in Figures 8–10 for the configurations at hours 94, 144, and 205 respectively. The panels shown are organized in the same order as the simple current sheet reconstruction. In Table 2 we show the volume (in units of  $10^6 \text{ km}^3$ ) of the magnetic field that can be reconstructed with errors less than 1, 5, and 10%. This is done for all three of the investigated spacecraft configurations, and using all five of the nine-spacecraft reconstruction methods,  $M_{1,1}$ ,  $M_{1,2}$ ,  $M_{1,3}$ ,  $M_{1,4}$ , and  $M_2$ . In the bottom half of this table, we compare the volume reconstructed using a single regular tetrahedron to that of our five reconstruction methods.

Near the barycenter of each of the nine-spacecraft configurations, located at the origin of Figures 8–10, the magnetic field can be reconstructed to within 1% accuracy. The second-order method,  $M_2$ , can only reconstruct the magnetic field to within 10% accuracy in a small region near the barycenter of the configuration, while the first-order methods can reconstruct the magnetic field within 10% over a much greater area. This is the case because the second-order Taylor series expansion diverges quadratically with distance away from the barycenter of the spacecraft configuration, while the first-order Taylor series only diverges linearly with distance. Since our goal is to maximize the volume of accurate reconstruction, the first-order methods are superior. However, the second-order method may be more accurate at reconstructing the values of



**TABLE 2 |** Volumes (in units of  $10^6 \text{ km}^3$ ) with reconstructed magnetic field error less than 1, 5, or 10% for the three configurations using the four first-order methods and the second-order method discussed in §2.2. These volumes are compared to the equivalent regions reconstructed from a single regular tetrahedron with the same characteristic size as the overall nine-spacecraft configuration.

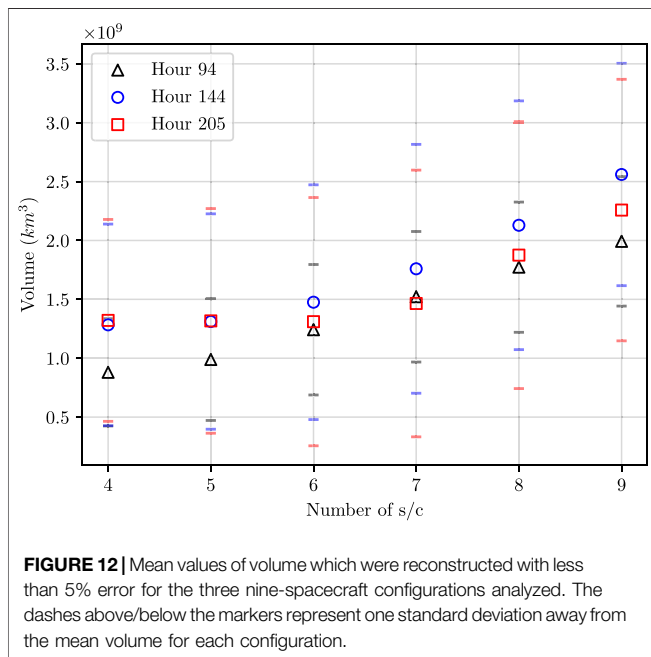
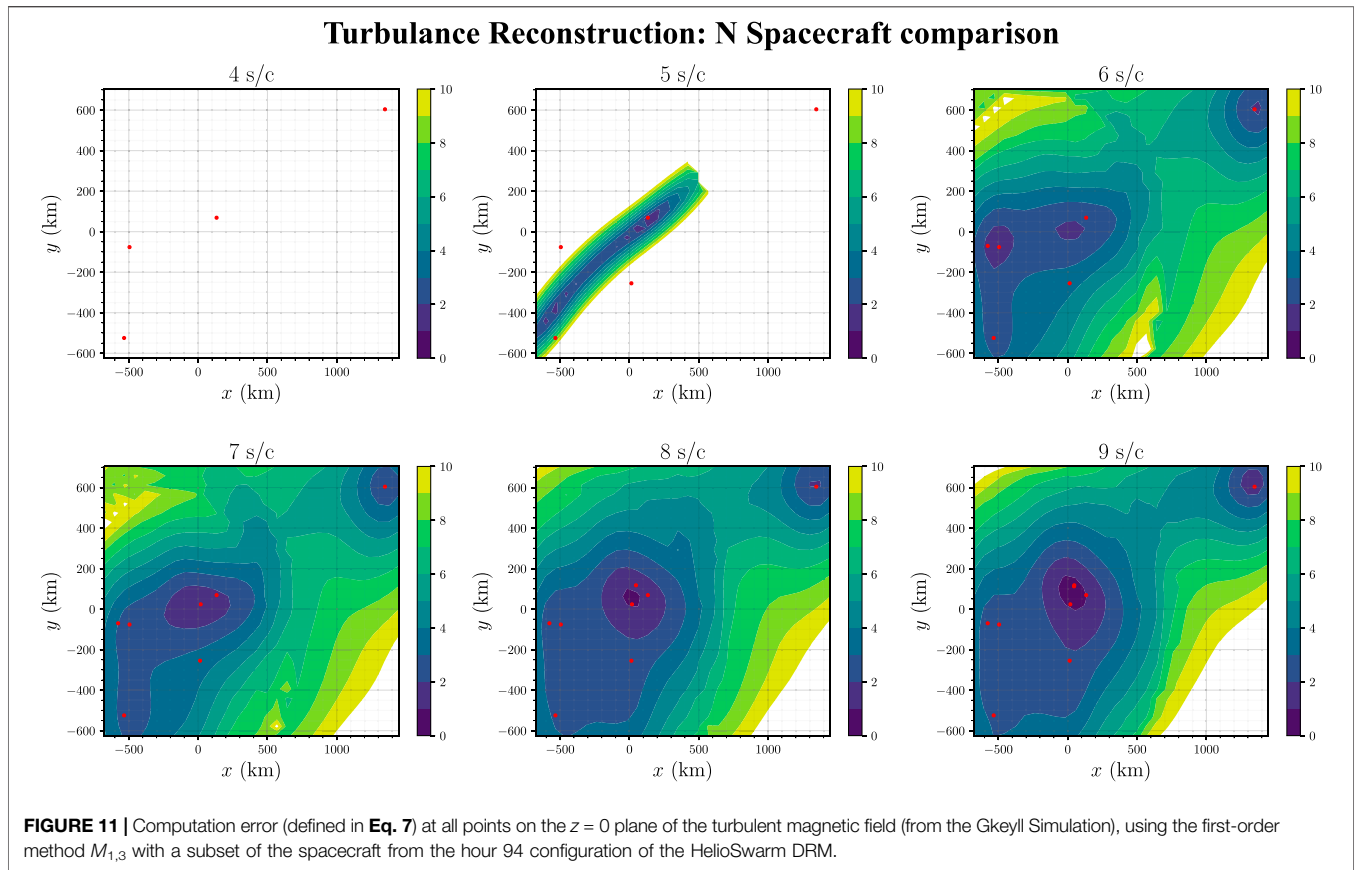
Volume	Hour 94			Hour 144			Hour 205		
	$\epsilon(1)$	$\epsilon(5)$	$\epsilon(10)$	$\epsilon(1)$	$\epsilon(5)$	$\epsilon(10)$	$\epsilon(1)$	$\epsilon(5)$	$\epsilon(10)$
$M_{1,1}$	17.56	1,057	3,095	2.95	324.4	2021	1.325	145.7	1,231
$M_{1,2}$	17.93	1,197	2,425	3.686	387.8	2,330	1.325	218.6	1,047
$M_{1,3}$	21.30	1816	3,208	9.584	2,151	4,722	12.59	1,556	4,334
$M_{1,4}$	30.26	1,189	1,679	17.69	1,169	2,281	0.0	0.0	0.0
$M_2$	3.363	47.45	118.1	5.898	48.66	137.9	3.975	33.79	103.3
Vs Regular Tetrahedron (%)	$\epsilon(1)$	$\epsilon(5)$	$\epsilon(10)$	$\epsilon(1)$	$\epsilon(5)$	$\epsilon(10)$	$\epsilon(1)$	$\epsilon(5)$	$\epsilon(10)$
$M_{1,1}$	137.20	54.31	38.70	24.58	23.75	25.26	13.80	10.77	15.39
$M_{1,2}$	140.12	61.51	30.32	30.72	28.40	29.13	13.80	16.16	13.08
$M_{1,3}$	166.39	93.33	40.12	79.87	157.54	59.04	131.12	115.03	54.18
$M_{1,4}$	236.44	61.06	21.00	147.45	85.62	28.51	0.0	0.0	0.0
$M_2$	26.27	2.44	1.48	49.15	3.56	1.72	41.41	2.50	1.29

the magnetic field very close to the barycenter of a spacecraft configuration.

The largest disparity compared to the current sheet simulations occurs for the single regular tetrahedron case, on the bottom right panels of each figure. The magnetic field is again only reconstructed accurately near each of the spacecraft, but because the turbulence simulation lacks angular symmetry, these regions manifest as spheres centered around each spacecraft. These four spheres appear to be the same size on the bottom

right panel of each figure because the spacecraft are equidistant from the  $z = 0$  plane.

Shown in Table 2, the single regular tetrahedron reconstructs the largest volume with less than 10% error, however the first-order methods reconstruct larger volumes with smaller errors. To maximize the volume reconstructed with less than 1% error, it appears it is best to use the first-order method  $M_{1,4}$ , detailed in §2.2.4 (if a sufficient number of quasi-regular tetrahedra can be formed from the nine spacecraft configuration).



**TABLE 3** | From the nine-spacecraft configurations of hours 94, 144, and 205, we select a subset of  $N \in \{4, 5, 6, 7, 8, 9\}$  spacecraft. We determine the probability that this  $N$  spacecraft configuration does not contain a tetrahedron which passes the threshold shape requirements of first-order reconstruction method  $M_{1,3}$ .

N	4	5	6	7	8	9
Hour 94	57.1	13.5	1.2	0	0	0
Hour 144	71.4	36.5	9.5	0	0	0
Hour 205	77.8	51.6	20.2	0	0	0

completed using the Monte Carlo sampling of the turbulent simulation as described in §3.

For  $N \in \{4, 5, 6, 7, 8, 9\}$  spacecraft, we reconstructed the value of the magnetic field at all  $30 \times 30 \times 30$  points  $\xi$  using all  $C(N, 4)$  tetrahedra. We then use the  $M_{1,3}$  first-order reconstruction method of §2.2.4 to reconstruct  $\mathbf{B}$  at all points  $\xi$ . The errors everywhere are computed using Eq. 7, and the volume where the error is less than 5% is computed using Eq. 8 multiplied by the total reconstructed volume. We visualize the errors of this method for the hour 94 configuration in Figure 11. In this example, we find that as the number of spacecraft is increased, the area which is reconstructed with a high accuracy also increases. As this result depends on which particular subset of spacecraft are chosen for a given  $N$ , we next investigate whether this increase is holds for an arbitrary selection of spacecraft.

### 3.3 Sensitivity to Number of Spacecraft

We analyze how the volume reconstructed with less than 5% error varies as a function of the number of spacecraft. This analysis was

We start by choosing four out of the 9 spacecraft of the hour 94 configuration. These spacecraft measurements are used to estimate the value of  $\mathbf{B}$  everywhere via the first-order reconstruction method  $M_{1,3}$ . We find the volume over which we can reconstruct  $\mathbf{B}$  with an error less than 5%. This process is repeated for all 126 possible choices of four spacecraft. We repeat all of these volume calculations, initializing the spacecraft configuration at 50 different locations within the simulated turbulent  $\mathbf{B}$  field. Finally, we take the mean of all  $126 \times 50$  volume values and plot them in **Figure 12**. In these averages, we omit the instances where no tetrahedra pass the selection criteria of method  $M_{1,3}$ . We repeat this process for  $N = 5, 6, 7, 8,$  and  $9$  spacecraft from the hour 94 configuration, as well as for the hour 144 and 205 configurations.

As shown in **Figure 12**, we see that increasing the number of spacecraft measurements available increases the volume of the magnetic field reconstructed with less than 5% error. The variance of this reconstructed volume is smallest for the hour 94 configuration, which contains the most tetrahedra which are quasi-regular ( $\chi_j \leq 1$ ). However, it is not the case that the hour 94 configuration has the highest average volume which is reconstructed with less than 5% error.

We also track the instances where zero of the available tetrahedra in the set of  $N$  spacecraft meet the shape threshold of  $\chi_j \leq 1$  for the  $M_{1,3}$  method. The percentage of arrangements where this occurs is shown in **Table 3** as a function of spacecraft configuration (hour) and number of spacecraft,  $N$ . We see from this table that for the analyzed configurations, there must be at least seven spacecraft measurements to guarantee that at least one tetrahedron passes the previously stated shape criteria.

## 4 DISCUSSION

We have demonstrated that our reconstruction methods are an effective way to leverage magnetometer measurements from a configuration consisting of more than four spacecraft. We have defined a shape metric,  $\chi$ , for a tetrahedron of spacecraft which can be used as a threshold criterion. Estimates of magnetic field derived from tetrahedron which do meet the threshold value of  $\chi$  will be discarded, as they are misshapen and therefore more likely to produce erroneous estimates. Finally, we have shown that increasing the number of spacecraft in a configuration will increase the volume over which the magnetic field can be accurately reconstructed, as well as increase the likelihood that some tetrahedra of spacecraft in the configuration are well shaped.

In **Table 2** we demonstrated that our second-order reconstruction method  $M_2$  does not reconstruct the magnetic field with high accuracy over a large volume. However, we have shown that methods  $M_{1,3}$  and  $M_{1,4}$ , which average over a subset of the many available tetrahedra formed by nine spacecraft, improves the field reconstruction. This work indicates that the subset of tetrahedra which should be averaged over needs to consider each tetrahedron's spacial proximity to the reconstructed point as well as its geometric properties. By comparing results from spacecraft configurations with different tetrahedral geometric configurations, we find that designing spacecraft trajectories which maximize the number of tetrahedra that are quasi-regular (i.e.,  $\chi \leq 1$ ) is essential to improving the accuracy of the reconstructed magnetic field.

This work can help optimize future multi-spacecraft missions, such as HelioSwarm. The selection of tetrahedra which are included in the calculation of  $\mathbf{B}$  can be tuned to maximize the volume over which the field is reconstructed accurately, or it can be tuned to recreate  $\mathbf{B}$  as accurately as possible over a small volume. The first-order methods discussed here can be applied to reconstruct any vector field which is sparsely sampled by *in-situ* measurements, as no assumptions are made about the physical properties of the field.

The first-order reconstruction method applied to a single tetrahedron reconstructs the magnetic field perfectly at each spacecraft location. However, using any of our proposed composite first-order reconstruction methods, which average over many of these reconstructions, negates this behavior. In future work, we plan to construct a weight function which, when introduced into the tetrahedral averaging, returns this desired limiting behavior. Additional future work could include characterizing methods of predicting the surface inside-of-which we have less than a prescribed error value for an arbitrary configuration of spacecraft.

The authors would like to thank the HelioSwarm Science and Flight Dynamics teams for discussions and comments during the execution of this project, in particular Laura Plice and Jonathan Niehof. This material is based upon High Performance Computing (HPC) resources supported by the University of Arizona TRIF, UITS, and Research, Innovation, and Impact (RII) and maintained by the UArizona Research Technologies department. The authors acknowledge the Texas Advanced Computing Center (TACC) at The University of Texas at Austin for providing HPC resources that have contributed to the research results reported within this paper. D.V. is supported by the Science and Technology Facilities Council (STFC) Ernest Rutherford Fellowship ST/P003826/1 and STFC Consolidated Grant ST/S000240/1. J.M.T is supported by NSF STR award AGS-1842638.

## DATA AVAILABILITY STATEMENT

The raw data supporting the conclusions of this article will be made available by the authors, without undue reservation.

## AUTHOR CONTRIBUTIONS

TB conducted the numerical experiments, analyzed the results, and wrote the majority of the paper, with guidance and oversight from KK on all steps. JT supplied the turbulence simulation results and wrote up its description. ID provided background work on second order methods which served as a starting point for the one we developed. OR and DV provided knowledge of previous missions and context to the results.

## SUPPLEMENTARY MATERIAL

The Supplementary Material for this article can be found online at: <https://www.frontiersin.org/articles/10.3389/fspas.2021.727076/full#supplementary-material>

## REFERENCES

- Angelopoulos, V. (2008). The THEMIS Mission. *Space Sci. Rev.* 141, 5–34. doi:10.1007/s11214-008-9336-1
- Burch, J. L., Moore, T. E., Torbert, R. B., and Giles, B. L. (2016). Magnetospheric Multiscale Overview and Science Objectives. *Space Sci. Rev.* 199, 5–21. doi:10.1007/s11214-015-0164-9
- Constantinescu, O. D., Glassmeier, K.-H., Motschmann, U., Treumann, R. A., Fornaçon, K.-H., and Fränz, M. (2006). Plasma Wave Source Location Using CLUSTER as a Spherical Wave Telescope. *J. Geophys. Res.* 111, A09221. doi:10.1029/2005JA011550
- Dai, L., Wang, C., Cai, Z., Gonzalez, W., Hesse, M., Escoubet, P., et al. (2020). AME: A Cross-Scale Constellation of CubeSats to Explore Magnetic Reconnection in the Solar–Terrestrial Relation. *Front. Phys.* 8, 89. doi:10.3389/fphy.2020.00089
- Dunlop, M. W., Southwood, D. J., Glassmeier, K. H., and Neubauer, F. M. (1988). Analysis of Multipoint Magnetometer Data. *Adv. Space Res.* 8, 273. doi:10.1016/0273-1177(88)90141-X
- Elsasser, W. M. (1950). The Hydromagnetic Equations. *Phys. Rev.* 79, 183. doi:10.1103/PhysRev.79.183
- Escoubet, C. P., Fehringer, M., and Goldstein, M. (2001). Introduction The Cluster mission. *Ann. Geophys.* 19, 1197–1200. doi:10.5194/angeo-19-1197-2001
- Forsyth, C., Lester, M., Fazakerley, A. N., Owen, C. J., and Walsh, A. P. (2011). On the Effect of Line Current Width and Relative Position on the Multi-Spacecraft Curlometer Technique. *Planet. Space ScienceSpace Sci.* 59, 598–605. doi:10.1016/j.pss.2009.12.007
- Fu, H. S., Vaivads, A., Khotyaintsev, Y. V., Olshevsky, V., André, M., Cao, J. B., et al. (2015). How to Find Magnetic Nulls and Reconstruct Field Topology with MMS Data? *J. Geophys. Res. Space Phys.* 120, 3758–3782. doi:10.1002/2015JA021082
- Fu, H. S., Wang, Z., Zong, Q., Chen, X. H., He, J. S., Vaivads, A., et al. (2020). Methods for Finding Magnetic Nulls and Reconstructing Field Topology. *Am. Geophys. Union.*, 153–172. doi:10.1002/9781119509592.ch9
- Goldreich, P., and Sridhar, S. (1995). Toward a Theory of Interstellar Turbulence. 2: Strong Alfvénic Turbulence. *Astrophysical J.* 438, 763. doi:10.1086/175121
- ISSI book (2008). “Multi-Spacecraft Analysis Methods Revisited,” in *ISSI Scientific Report SR-008*. Editors G. Paschmann and P. W. Daly (Bern, Switzerland: International Space Science Institute).
- Hakim, A., Loverich, J., and Shumlak, U. (2006). A High Resolution Wave Propagation Scheme for Ideal Two-Fluid Plasma Equations. *J. Comput. Phys.* 219, 418–442. doi:10.1016/j.jcp.2006.03.036
- Klein, K. G., Alexandrova, O., Bookbinder, J., Caprioli, D., Case, A. W., Chandran, B. D. G., et al. (2019). Plasma 2020 Decadal] Multipoint Measurements of the Solar Wind: A Proposed Advance for Studying Magnetized Turbulence. arXiv:1903.05740. Available at: <https://arxiv.org/abs/1903.05740>.
- Li, T. C., Howes, G. G., Klein, K. G., and TenBarge, J. M. (2016). Energy Dissipation and Landau Damping in Two- and Three-Dimensional Plasma Turbulence. *Astrophysical J.* 832, L24. doi:10.3847/2041-8205/832/2/L24
- Narita, Y., Nakamura, R., and Baumjohann, W. (2013). Cluster as Current Sheet Surveyor in the Magnetotail. *Ann. Geophys.* 31, 1605–1610. doi:10.5194/angeo-31-1605-2013
- Orszag, S. A., and Tang, C.-M. (1979). Small-Scale Structure of Two-Dimensional Magnetohydrodynamic Turbulence. *J. Fluid Mech.* 90, 129–143. doi:10.1017/S002211207900210X
- Paschmann, G., and Daly, P. W. (1998). *Analysis Methods for Multi-Spacecraft Data*. Netherlands: The International Space Science Institute, Hallerstr. 6, CH-3012 Bern, Switzerland. ESA Publications Division, Noordwijk. Copyright: 1998 ISSI/ESA. ISSI Scientific Report SR-001.
- Pinçon, J.-L., and Motschmann, U. (1998). “Multi-spacecraft Filtering: General Framework,” in *Analysis Methods for Multi-Spacecraft Data, ISSI Scientific Reports Series*. Editors G. Paschmann and P. W. Daly, 1, 65.
- Plice, L., Perez, A. D., and West, S. (2020). Advances in the Astronomical Sciences, 171.
- Robert, P., Dunlop, M. W., Roux, A., and Chanteur, G. (1998). Tetrahedron Geometric Factors. *ISSI Scientific Rep. Ser.* 1, 323–348.
- Schwartz, S., Bale, S. D., Fujimoto, M., Hellinger, P., Kessel, M., Le, G., et al. (2009). Cross-Scale: Multi-Scale Coupling in Space Plasma, Assessment Study Report. arXiv:0912.0856. Available at: <https://arxiv.org/abs/0912.0856>.
- Srinivasan, B., and Shumlak, U. (2011). Analytical and Computational Study of the Ideal Full Two-Fluid Plasma Model and Asymptotic Approximations for Hall-Magnetohydrodynamics. *Phys. Plasmas.* 18, 092113. doi:10.1063/1.3640811
- Torbert, R. B., Dors, I., Argall, M. R., Genestreti, K. J., Burch, J. L., Farrugia, C. J., et al. (2020). A New Method of 3-D Magnetic Field Reconstruction. *Geophys. Res. Lett.* 47, e85542. doi:10.1029/2019GL085542
- Verscharen, D., Klein, K. G., and Maruca, B. A. (2019). The Multi-Scale Nature of the Solar Wind. *Living Rev. Sol. Phys.* 16, 5. doi:10.1007/s41116-019-0021-0
- Wang, L., Hakim, A. H., Bhattacharjee, A., and Germaschewski, K. (2015). Comparison of Multi-Fluid Moment Models with Particle-In-Cell Simulations of Collisionless Magnetic Reconnection. *Phys. Plasmas.* 22, 012108. doi:10.1063/1.4906063
- Wang, L., Hakim, A. H., Ng, J., Dong, C., and Germaschewski, K. (2020). Exact and Locally Implicit Source Term Solvers for Multifluid-Maxwell Systems. *J. Comput. Phys.* 415, 109510. doi:10.1016/j.jcp.2020.109510

**Conflict of Interest:** The authors declare that the research was conducted in the absence of any commercial or financial relationships that could be construed as a potential conflict of interest.

**Publisher’s Note:** All claims expressed in this article are solely those of the authors and do not necessarily represent those of their affiliated organizations, or those of the publisher, the editors and the reviewers. Any product that may be evaluated in this article, or claim that may be made by its manufacturer, is not guaranteed or endorsed by the publisher.

Copyright © 2021 Broeren, Klein, TenBarge, Dors, Roberts and Verscharen. This is an open-access article distributed under the terms of the Creative Commons Attribution License (CC BY). The use, distribution or reproduction in other forums is permitted, provided the original author(s) and the copyright owner(s) are credited and that the original publication in this journal is cited, in accordance with accepted academic practice. No use, distribution or reproduction is permitted which does not comply with these terms.