



An Orientation-Based Unification of Young Jetted AGN: The Case of 3C 286

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In recent years, the old paradigm according to which only high-mass black holes can launch powerful relativistic jets in active galactic nuclei (AGN) has begun to crumble. The discovery of γ -rays coming from narrow-line Seyfert 1 galaxies (NLS1s), usually considered young and growing AGN harboring a central black hole with mass typically lower than $10^8 M_{\odot}$, indicated that also these low-mass AGN can produce powerful relativistic jets. The search for parent population of γ -ray emitting NLS1s revealed their connection with compact steep-spectrum sources (CSS). In this proceeding we present a review of the current knowledge of these sources, we present the new important case of 3C 286, classified here for the first time as NLS1, and we finally provide a tentative orientation based unification of NLS1s and CSS sources.

Keywords: Active Galactic Nuclei (AGN), narrow-line Seyfert 1 galaxies, compact steep-spectrum sources, relativistic jets, unification models

1. OVERVIEW

Relativistic jets are usually thought to be a product of accretion onto a central supermassive black hole (SMBH) (e.g., Blandford and Znajek, 1977; Blandford and Payne, 1982). In active galactic nuclei (AGN), the interplay between the accretion disk and the SMBH is likely to produce powerful and collimated bipolar outflows. According to the unified model of jetted-AGN (Urry and Padovani, 1995), the two classes of blazars, BL Lacertae objects (BL Lacs) and flat-spectrum radio quasars (FSRQs), are double-lobed radio galaxies, with Fanaroff-Riley (FR) morphological types I and II, respectively, in which the line of sight falls inside the relativistic jet aperture cone. In recent years, this scenario has been slightly revised in a more physical way. BL Lacs and FSRQs indeed seem to reflect two different accretion modes. While BL Lacs have an inefficient accretion onto their central SMBH and a low density environment, FSRQs have instead a strong disk accreting efficiently and a photon- and matter-rich environment (Best and Heckman, 2012). This dichotomy led to the distinction between high-excitation radio galaxies and low-excitation radio galaxies (HERGs and LERGs, respectively) classified according, for example, to the [O III]/H α flux ratio (Laing et al., 1994), and corresponding, respectively, to FSRQ and BL Lac objects when observed at small angles (although exceptions are known, Giommi et al., 2012, see also the review by Foschini, 2017).

Another important aspect of jetted-AGN is that for many years they were thought to be produced only by the most massive black holes harbored in giant passive elliptical galaxies (Laor, 2000; Böttcher and Dermer, 2002; Marscher, 2009). However, the discovery of gamma-ray emission coming from the AGN class known as narrow-line Seyfert 1 galaxies (NLS1s) proved this paradigm

to be wrong (Abdo et al., 2009a,b; Foschini et al., 2015). Low mass AGN, indeed, are also able to launch relativistic jets, whose power is lower because of the well-known scaling of the black hole mass (Heinz and Sunyaev, 2003; Foschini, 2011). This result was confirmed by the identification of low-mass sources among flat-spectrum radio quasars (FSRQs, Ghisellini, 2016), and by the discovery of relativistic jets associated with spiral galaxies (e.g., Keel et al., 2006; Mao et al., 2015).

NLS1s are the most prominent class of low-mass AGN. Classified according to the $\text{FWHM}(H\beta) < 2,000 \text{ km s}^{-1}$ and the ratio between $[\text{O III}]/H\beta < 3$, they are unobscured AGN with relatively low black hole mass ($\lesssim 10^8 M_{\odot}$) and high Eddington ratio, which accounts for the narrowness of the permitted lines and other observational properties (for a review, see Komossa, 2008). A well-known interpretation for these peculiar features is that NLS1s are young AGN still growing and evolving (Mathur, 2000), or sources rejuvenated by a recent merger (Mathur et al., 2012). While a large fraction of NLS1s is radio-quiet, some of them are radio-loud and, as previously mentioned, harbor powerful relativistic jets (Yuan et al., 2008). These radio-loud NLS1s (RLNLS1s) are characterized by a strong and very fast variability consistent with the low-mass scenario, and particularly evident in those with a flat radio spectrum (Foschini et al., 2015).

Another class of radio-loud AGN is known as compact steep-spectrum sources (CSS, see the review by O’Dea, 1998). The radio morphology of these objects is typically characterized by fully developed radio lobes and by a small linear size, with jets confined within the host galaxy. Many authors interpret these properties as signs of young age (Fanti et al., 1995): different estimates of this parameter indeed revealed that they could be radio sources younger than 10^5 years (Owsianik and Conway, 1998; Murgia et al., 1999). CSS sources might eventually evolve to giant double-lobed radio galaxies, or simply switch off because of disk instabilities, going through multiple activity phases during their lifetime (e.g., Czerny et al., 2009; Orienti, 2015). As regular radio galaxies, CSS sources can also be divided in HERGs and LERGs according to their optical spectrum (Kunert-Bajraszewska and Labiano, 2010; Kunert-Bajraszewska et al., 2010).

2. CSS SOURCES AND NLS1S

A possible link between NLS1s and CSS was first suggested by Oshlack et al. (2001), who remarked how the RLNLS1 PKS 2004-447 radio morphology was consistent with that of a CSS source. In the following years, many authors reached the same conclusion (Gallo et al., 2006; Komossa et al., 2006; Yuan et al., 2008; Caccianiga et al., 2014; Schulz et al., 2015). In particular Gu et al. (2015) carried out a VLBA survey on 14 NLS1s at 5 GHz, finding that essentially all RLNLS1s with a steep radio spectrum have a compact morphology, just like CSS. Caccianiga et al. (2017), focusing on the steep-spectrum RLNLS1 SDSS J143244.91+301435.3, found a spectral turnover at low frequency and a small linear size, confirming that this object can be classified as CSS as well.

Beside the morphological similarities, there are other hints that point toward a unification between CSS sources and RLNLS1s. This was investigated under three different points of view in Berton et al. (2016a), using two complete samples of CSS/HERGs and flat-spectrum RLNLS1s (F-NLS1s) limited to $z = 0.6$. A first indication of similarity between these two samples was provided by the V/V_{max} test, which revealed that both of them are consistent with no evolution up to their redshift limit. Moreover, the 1.4 GHz luminosity function (LF) of both samples was derived. Following Urry and Padovani (1995), the effect of relativistic beaming was applied to the observed LF of CSS/HERG to predict the expected shape of the LF of the sources observed along the jet direction. The resulting LF was in good agreement with the observed LF of F-NLS1 thus indicating that at least some CSS/HERGs might be RLNLS1s observed at large angles.

Finally, a comparison between the black hole mass distributions of both samples unveiled again a very good agreement. We repeated here the analysis on the black hole mass distributions on a larger sample of 60 NLS1s. In **Figure 1** we show the distribution of the CSS/HERGs sample by Berton et al. (2016a), compared to all of the masses derived for flat- and steep-spectrum RLNLS1s in Foschini et al. (2015) (42 objects) and Berton et al. (2015) (17 sources), and to that of 3C 286 (see Section 3). It is immediately evident from the plot that 3C 286 is located in the same region as NLS1s. To test the similarity of the distributions, we applied the Kolmogorov-Smirnov test (K-S). The null hypothesis is that the two distributions are originated by the same population. As a rejection threshold, we used a p -value lower than 0.10. Considering CSS/HERGs, the null hypothesis cannot be rejected when comparing them to both flat- and steep-spectrum RLNLS1s, with p -values 0.82 and 0.39, respectively. Testing RLNLS1s as a whole, the p -value is 0.68, again too high to reject the null hypothesis. The mass distribution of CSS/HERGs is therefore statistically consistent

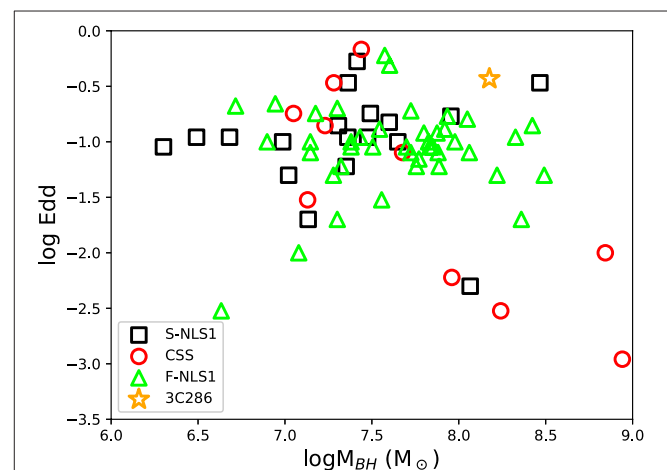


FIGURE 1 | Black hole mass vs. Eddington ratio distribution. Flat-spectrum RLNLS1s are indicated by green triangles, the red circles indicate CSS/HERG sources, the black squares indicate steep-spectrum RLNLS1s, and the orange star represents 3C 286.

with that of RLNLS1s, in agreement with the previously shown results.

3. 3C 286

We present here another interesting case of CSS/HERG, that of 3C 286 ($z = 0.85$). This object is a well-known CSS classified for the first time by Peacock and Wall (1982). It is often used as a calibrator at centimeter wavelengths for both total flux and linear polarization. It shows both the turnover in the radio spectrum around 300 MHz, a spectral index between 1.4 and 50 GHz of $\alpha = -0.61$. It also exhibits the compact morphology typical of CSS: while it remains unresolved at kpc scales, a core-jet structure is visible on pc scale. The jet inclination has been estimated of 48° with respect to the line of sight (e.g., see An et al., 2017, and references therein).

The first optical spectrum of this radio source was obtained by Schmidt (1962), who did not recognize the strongly redshifted Mg II $\lambda 2798$ line. Since its redshift determination (Burbidge and Burbidge, 1969), a significant improvement has been made in optical with the SDSS-BOSS survey (Dawson et al., 2013), which obtained a new spectrum extended up to 10000\AA . The $H\beta$ line, now visible, has a FWHM of $(1811 \pm 169) \text{ km s}^{-1}$, a ratio $[\text{O III}]/H\beta$ of 1.03 ± 0.05 , and Fe II multiplets are clearly visible in the spectrum (see Figure 2). The corresponding uncertainties were calculated with a Monte Carlo method described by Berton et al. (2016b). These parameters clearly indicate that 3C 286 can be classified as NLS1.

We derived its black hole mass following the method described by Foschini et al. (2015). After correcting for redshift, we subtracted the continuum emission by fitting it with a power

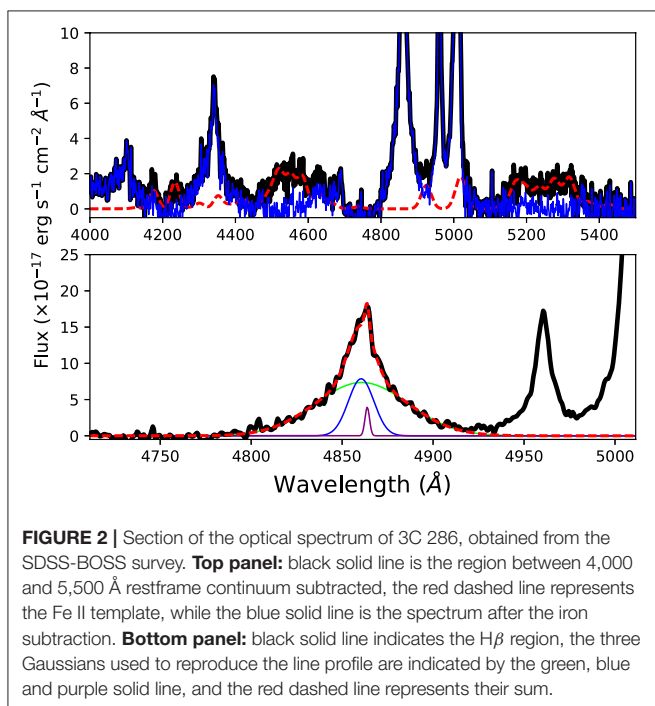


FIGURE 2 | Section of the optical spectrum of 3C 286, obtained from the SDSS-BOSS survey. **Top panel:** black solid line is the region between 4,000 and 5,500 Å restframe continuum subtracted, the red dashed line represents the Fe II template, while the blue solid line is the spectrum after the iron subtraction. **Bottom panel:** black solid line indicates the $H\beta$ region, the three Gaussians used to reproduce the line profile are indicated by the green, blue and purple solid line, and the red dashed line represents their sum.

law, plus the Fe II multiplets using an online software (Kovačević et al., 2010; Shapovalova et al., 2012)¹ as shown in top panel of Figure 2. We then fitted $H\beta$ profile using three Gaussians, one to reproduce the narrow component and two to reproduce the broad component (see bottom panel of Figure 1). The narrow component flux was fixed to be 1/10 of $[\text{O III}] \lambda 5007$ (Véron-Cetty et al., 2001) and its FWHM to be the same as $[\text{O II}] \lambda 3727$ ($\sim 600 \text{ km s}^{-1}$). After subtracting the $H\beta$ narrow component, we derived the second-order moment of the broad component σ as a proxy for velocity (Peterson et al., 2004), we obtained the broad-line region (BLR) radius by means of $H\beta$ luminosity. Under the hypothesis of virialized system, the black hole mass is $1.5 \times 10^8 M_\odot$ and the Eddington ratio, defined as the ratio between the bolometric luminosity and the Eddington luminosity of the black hole, is 0.37. Both of these values are consistent with those derived for other radio-loud NLS1s (Foschini et al., 2015, see Figure 1).

Interestingly, this object is a γ -ray emitter, identified in the 3rd Fermi catalog with the source 3FGL J1330.5+3023 (Ackermann et al., 2015)². 3C 286 is therefore the third γ -ray detected steep-spectrum RLNLS1 after RX J2314.9+2243 (Komossa et al., 2015) and B3 1441+476 (Liao et al., 2015). This is a very interesting result, since the number of misaligned NLS1s detected in γ -rays appears to be significantly higher than in other AGN classes. Indeed, only ~ 10 radiogalaxies are known to be γ -ray emitters (Rieger, 2016), and their corresponding beamed population of FSRQs and BL Lacs is made of 1144 sources. Conversely, among NLS1s, currently three misaligned γ -ray sources are known, out of nine oriented objects. The physical interpretation of this phenomenon is not clear yet, but interaction between the jet and the interstellar medium inside the host galaxy could in principle produce the observed γ -ray emission (Migliori et al., 2014).

The classification of this source has another important implication on NLS1s nature. Some authors (e.g., Decarli et al., 2008) suggested that NLS1s are not objects with a low-mass black hole and a high accretion rate. If their BLR had a disk-like shape, when observed pole-on the permitted lines would appear as narrow because of the lack of Doppler broadening. NLS1 would be just low-inclination sources rather AGN with a low-mass central SMBH. The large observing angle (48°) estimated by An et al. (2017) for 3C 286 seems to contradict this hypothesis, at least for this particular object and other RLNLS1s sharing the same property (e.g., Mrk 783, Congiu et al., 2017).

4. YOUNG JETTED-AGN UNIFICATION

As previously mentioned, CSS sources are usually considered as young radiogalaxies still growing (Fanti et al., 1995). The previous results seem to point out that these objects can often be jetted-NLS1s observed at large angles. Of course, this does not mean that all CSS sources are misaligned NLS1s, because some objects probably have a larger black hole mass. The opposite instead might be true, since many of the jetted-NLS1 observed at large angles studied so far appear as CSS/HERGs. An important conclusion can be inferred from this result. If CSS/HERGs are

¹http://servo.aob.rs/FeII_AGN/,

²<http://www.asdc.asi.it/fermi3fgl/>.

young sources, and RLNLS1s are CSS sources, they should be young objects as well. This result, just like the high inclination of 3C 286, is in contrast with the vision of NLS1s as a pole-on view of a type 1 AGN.

According to the unified model of radio-loud AGN (Urry and Padovani, 1995), the parent population of FSRQs are FR II radio galaxies, that is when an FR II is observed along its relativistic jet, it appears as a FSRQ. In a similar way, FR I radio galaxies constitute the parent population of BL Lacs. This interpretation might not be completely appropriate, since there are FR I sources that can be associated with FSRQs, and FR II which appear as BL Lacs. Recently it was suggested that a more accurate unification is between FSRQs and all the FR radio galaxies optically classified as HERG (FR_{HERG}), while BL Lacs should be associated to FR with a LERG-type optical spectrum (FR_{LERG}, Giommi et al., 2012). This revised unification is more physics-based, since it accounts for the different accretion mechanisms (strong disk for FSRQs, weak disk for BL Lacs, Best and Heckman, 2012). CSS/HERGs, in this picture, are the young stage of FR_{HERG} objects. Therefore, if flat-spectrum RLNLS1s are beamed CSS/HERGs, and FSRQs are beamed FR_{HERG}, F-NLS1s could be the young stage of FSRQs (Foschini et al., 2015; Berton et al., 2016a, see **Figure 3**).

This is in agreement with the finding that flat-spectrum RLNLS1s are the low-mass tail of FSRQs (Foschini et al., 2015; Paliya et al., 2016). Indeed, the jet power scales with the black hole mass (Heinz and Sunyaev, 2003; Foschini, 2011) and, as expected, the jet power of RLNLS1s is lower than that observed in FSRQs. Once rescaled for the central mass, though, the normalized power is essentially the same, indicating that the launching mechanism is the same (Foschini et al., 2015). Finally, as shown in Berton et al. (2016a), the radio luminosity function of flat-spectrum RLNLS1s suggests that they might be the low-luminosity tail of FSRQs, indicating again an evolutionary link between these two classes.

All these results can be summarized in the scheme of **Figure 3**. A young jetted-AGN with a strong accretion disk and a photon-rich environment in the central engine surroundings, when observed along its relativistic jets, can appear as a flat-spectrum RLNLS1, characterized by a low black hole mass and a low jet luminosity and power. When the inclination angle increases, the object appears first as a steep-spectrum RLNLS1, with an extended radio emission and a relatively small linear size, that can be classified as a CSS/HERG. Finally, when the line of sight intercepts the molecular torus surrounding the nucleus, the object appears as a type 2 AGN in optical and as a CSS in radio. For older jetted-AGN, instead, the usual unified model holds, with FR_{HERG} (type 1 or 2 quasars in optical, according to the inclination) as parent population of FSRQs.

Such a scheme is probably true only in a statistical sense. Objects like 3C 286 are still fairly rare, and might represent an exception. Moreover, it is unlikely that all type 1 (unobscured) young AGN fit the NLS1 classification. Several other non-NLS1 unobscured AGN with a low-mass black hole exist and were found among γ -ray emitting sources (e.g., Shaw et al., 2012; Ghisellini, 2016). These objects do not appear as NLS1s even when observed at larger angles and in fact, as expected, not all

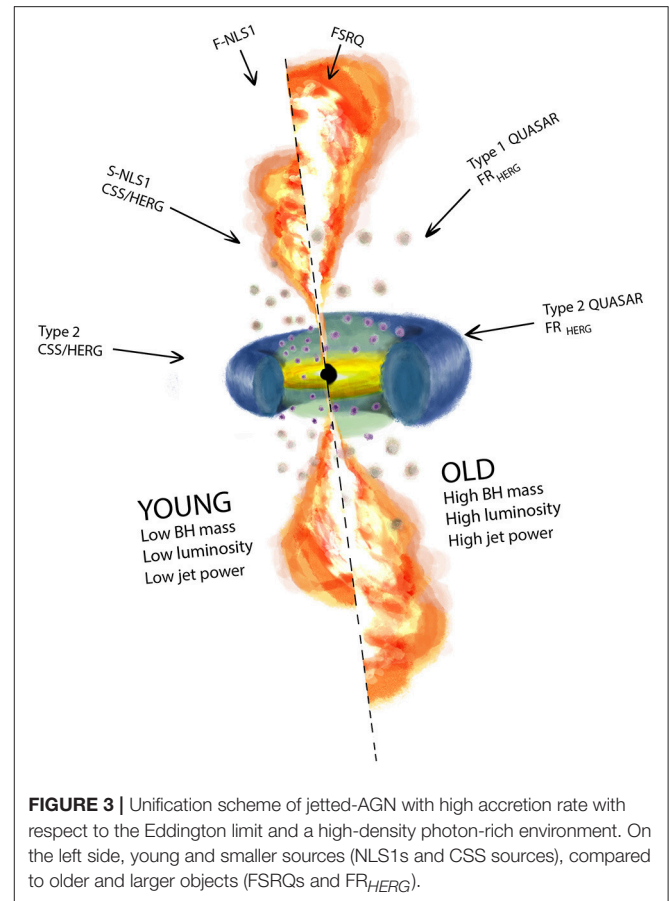


FIGURE 3 | Unification scheme of jetted-AGN with high accretion rate with respect to the Eddington limit and a high-density photon-rich environment. On the left side, young and smaller sources (NLS1s and CSS sources), compared to older and larger objects (FSRQs and FR_{HERG}).

type 1 CSS/HERGs can be classified as NLS1s. However, despite these notable exceptions, the scheme of **Figure 3** seems to be consistent with several observations.

This orientation-based unification could be confirmed by means of RLNLS1s and CSS/HERGs host galaxy observations. NLS1s usually have a late-type host galaxy (Crenshaw et al., 2003), and some RLNLS1s indeed seem to share the same property (Antón et al., 2008; Caccianiga et al., 2015; Kotilainen et al., 2016; Olguín-Iglesias et al., 2017). CSS sources in general are believed to be instead hosted by early-type hosts (e.g., Best et al., 2005). However, some studies pointed out that their host galaxy can be a spiral as well (Morganti et al., 2011), and that HERGs and LERGs can live in different morphological types, with HERGs showing more late-type hosts (Best and Heckman, 2012). Further observations are needed to better investigate this crucial aspect of young AGN.

5. SUMMARY

In this proceeding we presented a review of the current knowledge about young jetted-AGN, showing as well new results strictly connected to this topic and a tentative orientation-based unification model of these sources. We provided a new classification for the well-known CSS source 3C 286 which, thanks to new observations in the near-infrared, can be now

classified as a NLS1 galaxy. This result confirms that RLNLS1s can often appear, in radio, as CSS sources. Moreover, we confirmed that the black hole mass distribution of CSS/HERGs is very similar to that of RLNLS1s. Along with other results already published in the literature, these findings seem to strengthen the scheme according to which flat-spectrum RLNLS1s are CSS/HERGs observed along their relativistic jet, that can appear as a steep-spectrum RLNLS1 when observed at intermediate angles, and as a type 2 radio galaxy when obscured. RLNLS1s and CSS/HERGs then might represent the young and growing stage of FSRQs and FR_{HERG} radio galaxies, respectively, and constitute a part of the young AGN population. New observations on larger samples, particularly aimed at determining the host galaxy morphology, are needed to provide a final confirmation to this unification model.

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AUTHOR CONTRIBUTIONS

MB coordinated the project, analyzed the spectroscopic data, contributed to the development of the model, and wrote the manuscript. LF initiated the project, contributed to the development of the model, and provided a critical review of the manuscript. AC contributed to the development of the model, and provided a critical review of the manuscript. SC, EC, VC,

MF, GL, and PR helped collecting the data and provided a critical review of the manuscript.

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